Galaktikus archeológia

Kiss László, MTA CSFK CSI

Bevezetés a csillagászatba 4.

Miről lesz szó?

- A nagy vöröseltolódású Univerzum a közelben
- A Tejútrendszer szerkezete
- Égboltfelmérő programok: képek, színek, színképek, távolságok
- Módszerek a múlt rekonstruálására

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EVIDENCE FROM THE MOTIONS OF OLD STARS THAT THE GALAXY COLLAPSED

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Mount Wilson and Palomar Observatories
Carnegie Institution of Washington, California Institute of Technology
Received May 17, 1962

ABSTRACT

The (U, V, W)-velocity vectors for 221 well-observed dwarf stars have been used to compute the eccentricities and angular momenta of the galactic orbits in a model galaxy. It is shown that the eccentricity and the observed ultraviolet excess are strongly correlated. The stars with the largest excess (i.e., lowest metal abundance) are invariably moving in highly elliptical orbits, whereas stars with little or no excess move in nearly circular orbits. Correlations also exist between the ultraviolet excess and the W-velocity. Finally, the excess and the angular momentum are correlated; stars with large ultraviolet excesses have small angular momenta.

These correlations are discussed in terms of the dynamics of a collapsing galaxy. The data require that the oldest stars were formed out of gas falling toward the galactic center in the radial direction and collapsing from the halo onto the plane. The collapse was very rapid and only a few times 10^8 years were required for the gas to attain circular orbits in equilibrium (i e., gravitational attraction balanced by centrifugal acceleration). The scale of the collapse is tentatively estimated to be at least 10 in the radial direction and 25 in the Z-direction. The initial contraction must have begun near the time of formation of the first stars, some 10^{10} years ago.

I. INTRODUCTION

Beginning with Strömberg's early investigations of stellar motions, it has become

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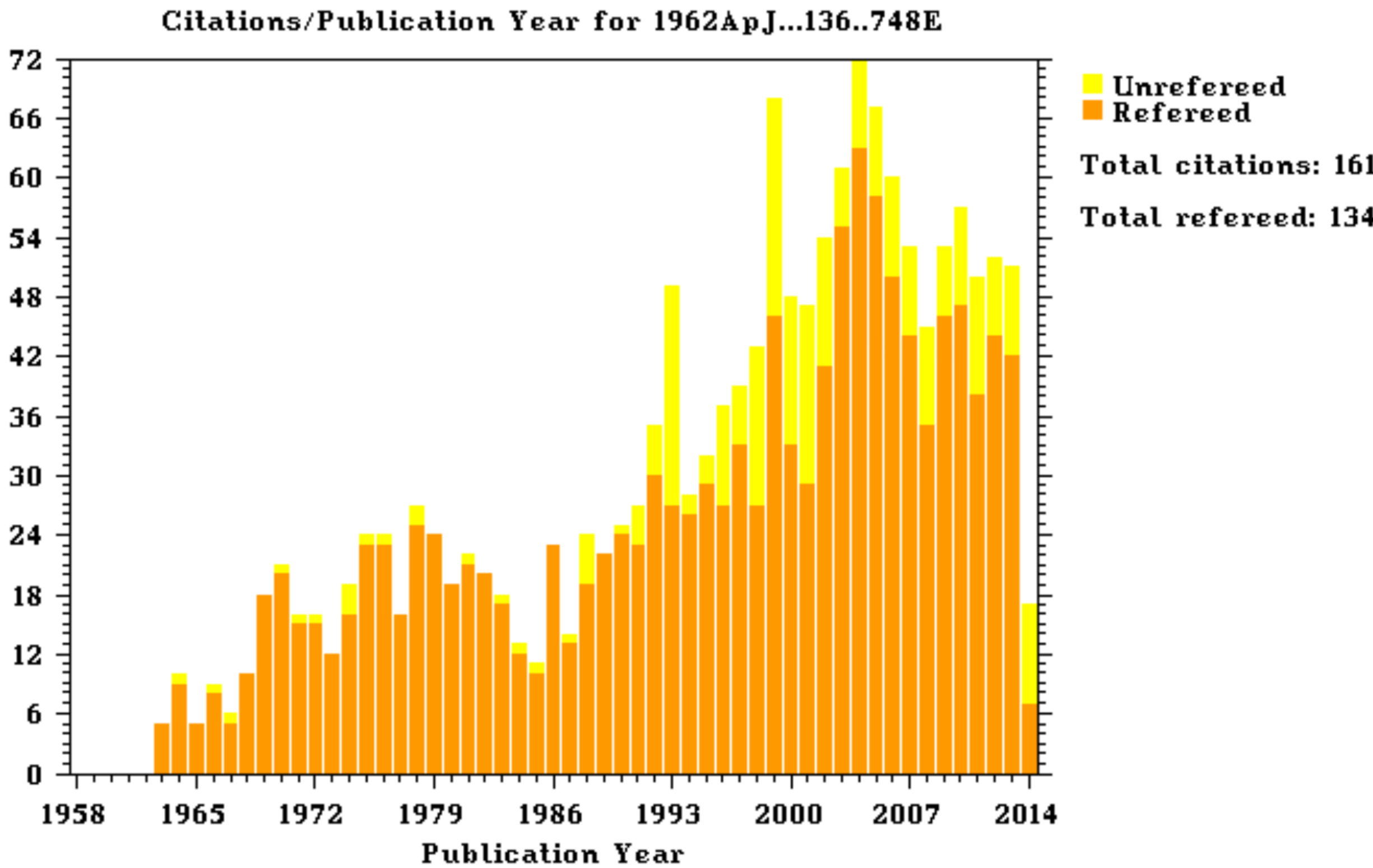
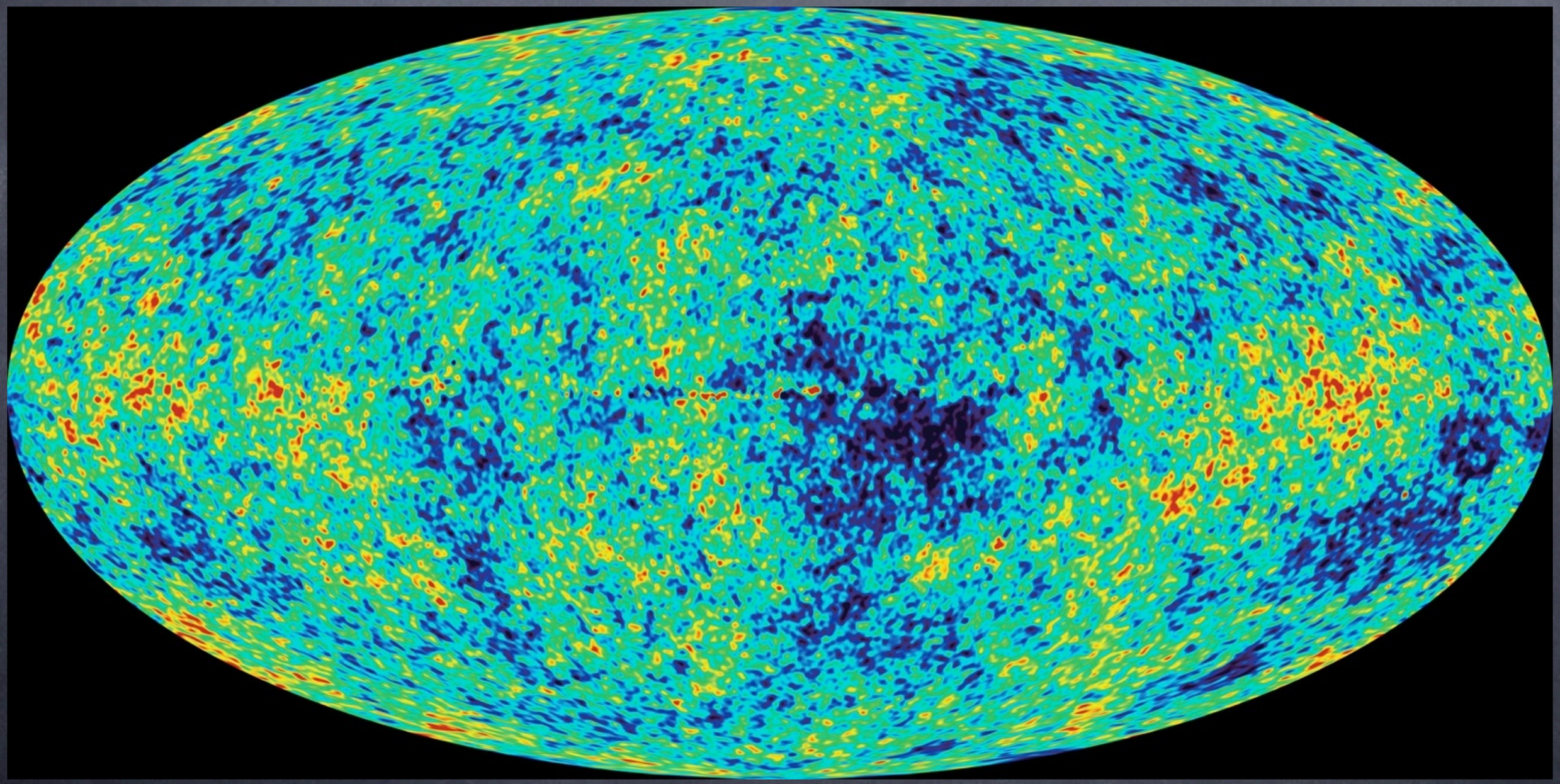


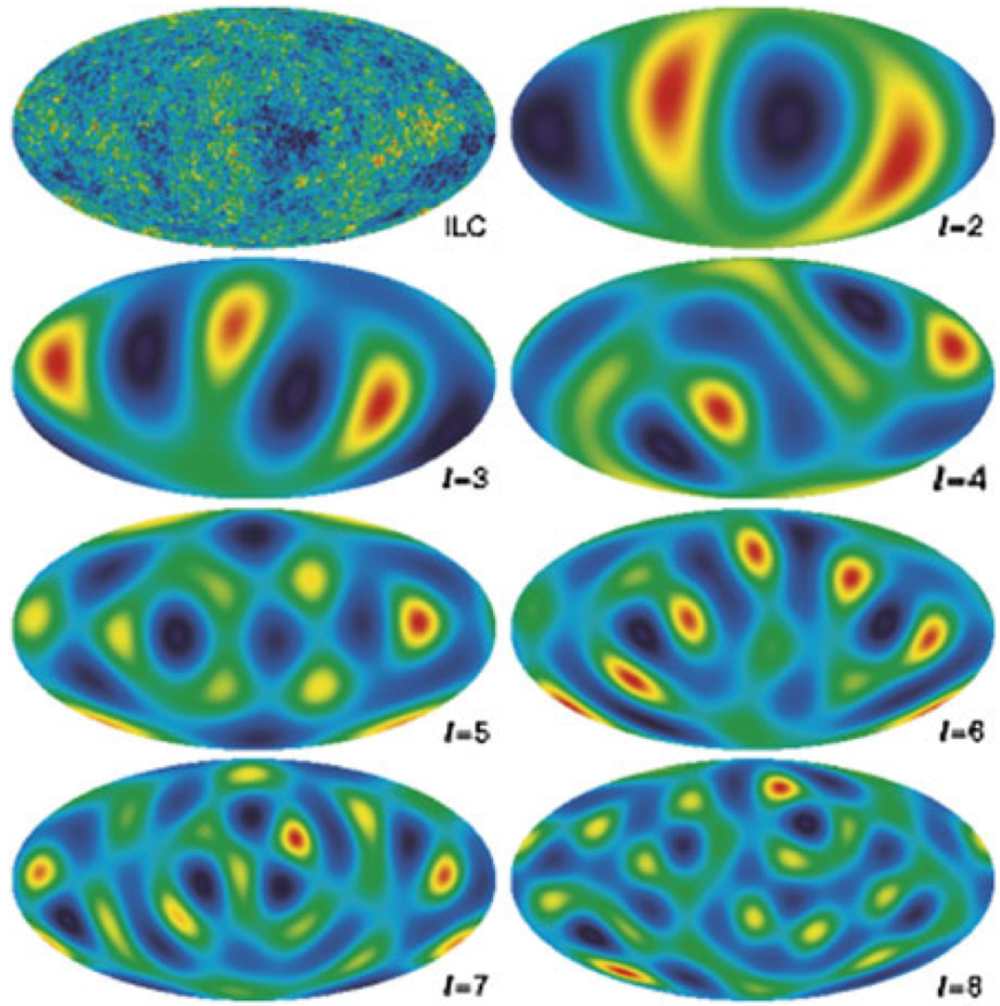
Chapter 1
Near Field Cosmology: The Origin of the Galaxy and the Local Group

Joss Bland-Hawthorn and Kenneth Freeman

The Galaxy has built up through a process of accretion and merging over billions of years which continues to this day. Astronomers are now embarking on a new era of massive stellar surveys over the coming decade. These campaigns will derive three-dimensional space motions and heavy element abundances for millions of stars throughout the Galaxy and its neighbours. The new observations will reveal signatures of the formation and early evolution of the Local Group; this is what we mean by 'near field cosmology.' We set this new course of study within the context of fossil signatures from galaxy surveys and the high redshift universe. We discuss the complex relationship between baryons and dark matter over cosmic time, and introduce a synthetic framework that will allow both numerical simulations and the impending data deluge to be compared. We also include relevant source materials for the young near-field cosmologist and some historical perspectives.

A WMAP térképe a mikrohullámú háttérsugárzásról: csomósodás irány szerint

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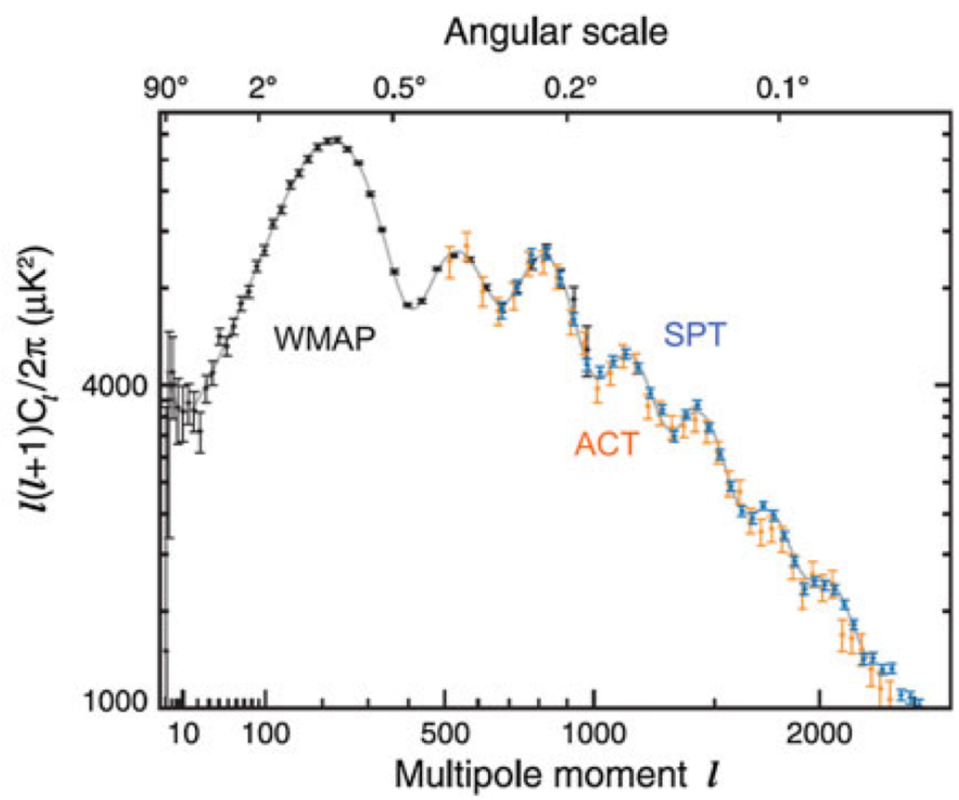
Angular scale
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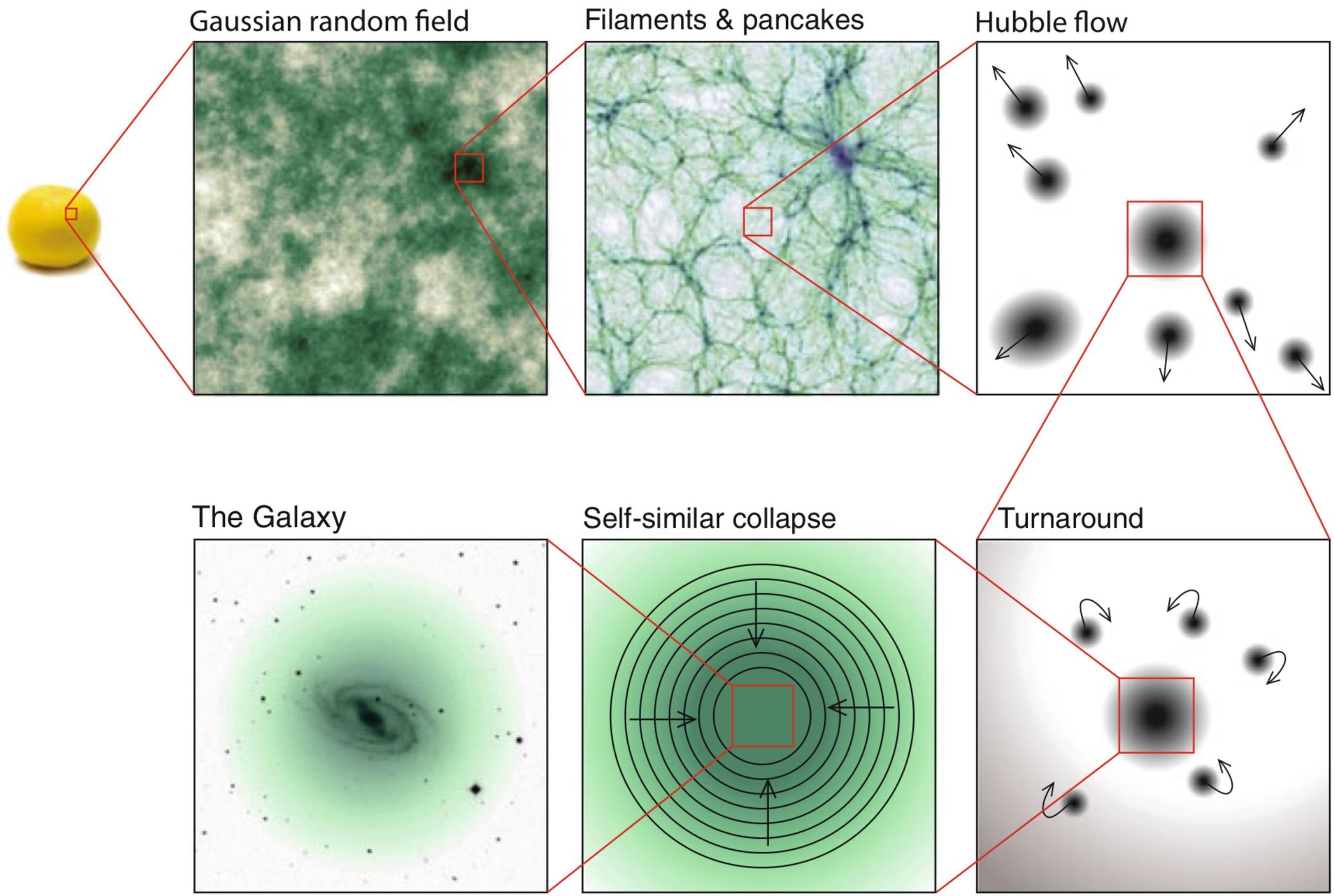


Fig. 1.15 A sequence of images depicting the major stages in the evolution of the Galaxy from the initial fluctuation spectrum—see text. This was preceded by inflation when the Universe went from sub-atomic scales to the size of a grapefruit in less than a picosecond

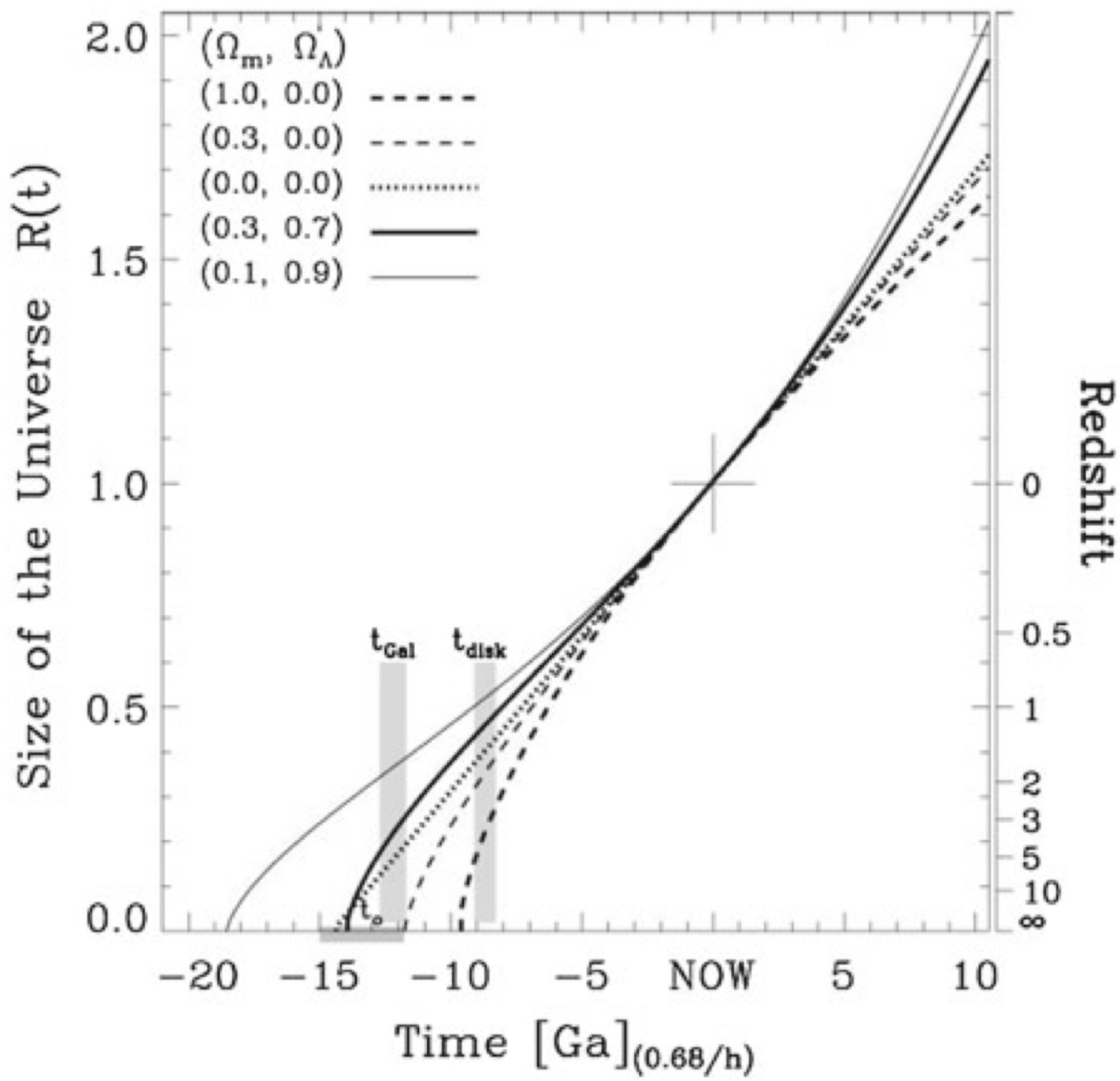
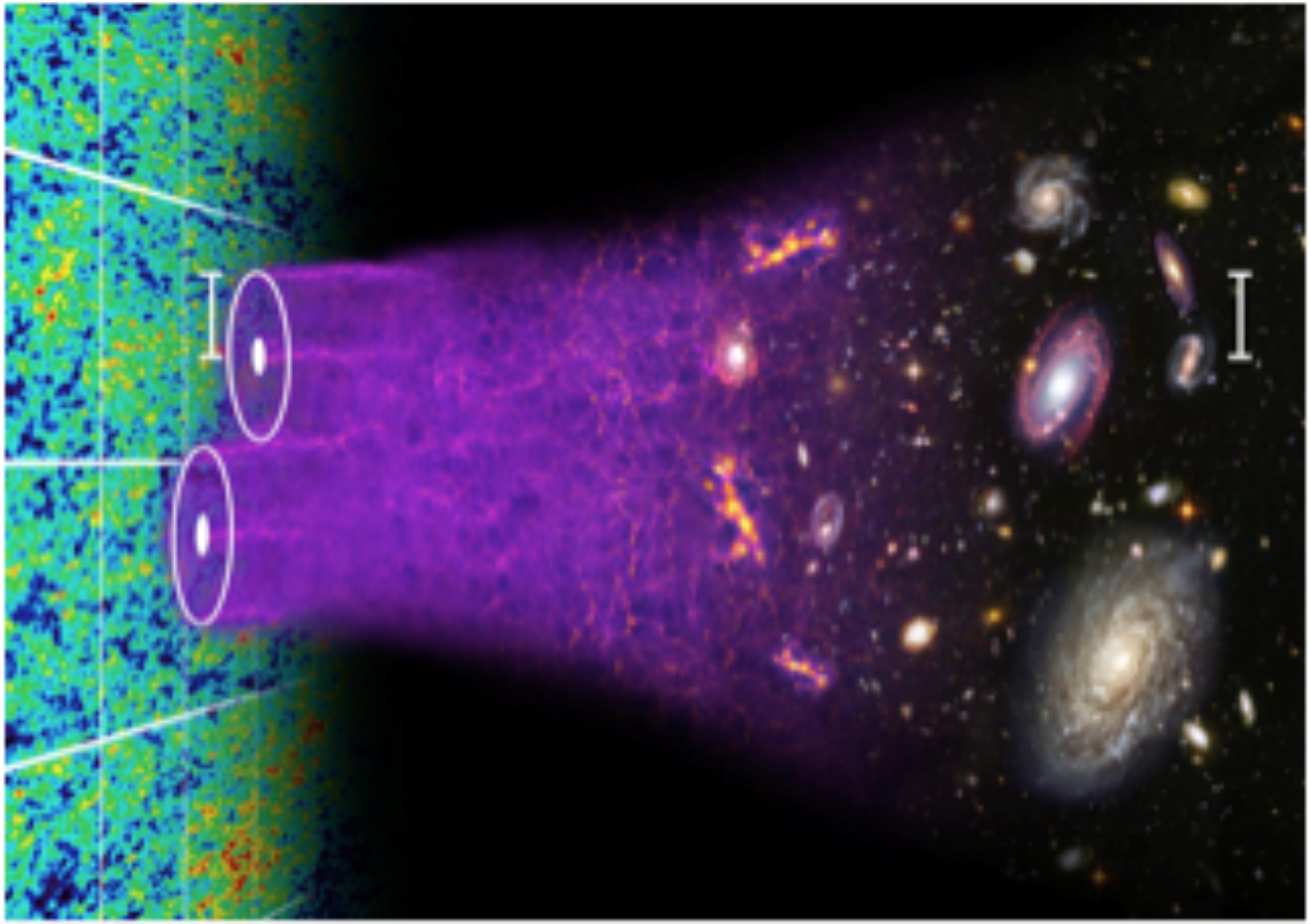


Fig. 1.14 Look-back time as a function of redshift and the size of the Universe for five different world models. The approximate ages of the Galactic halo and disk are indicated by hatched regions (we thank C. Lineweaver for modifying an earlier version of this figure for us)



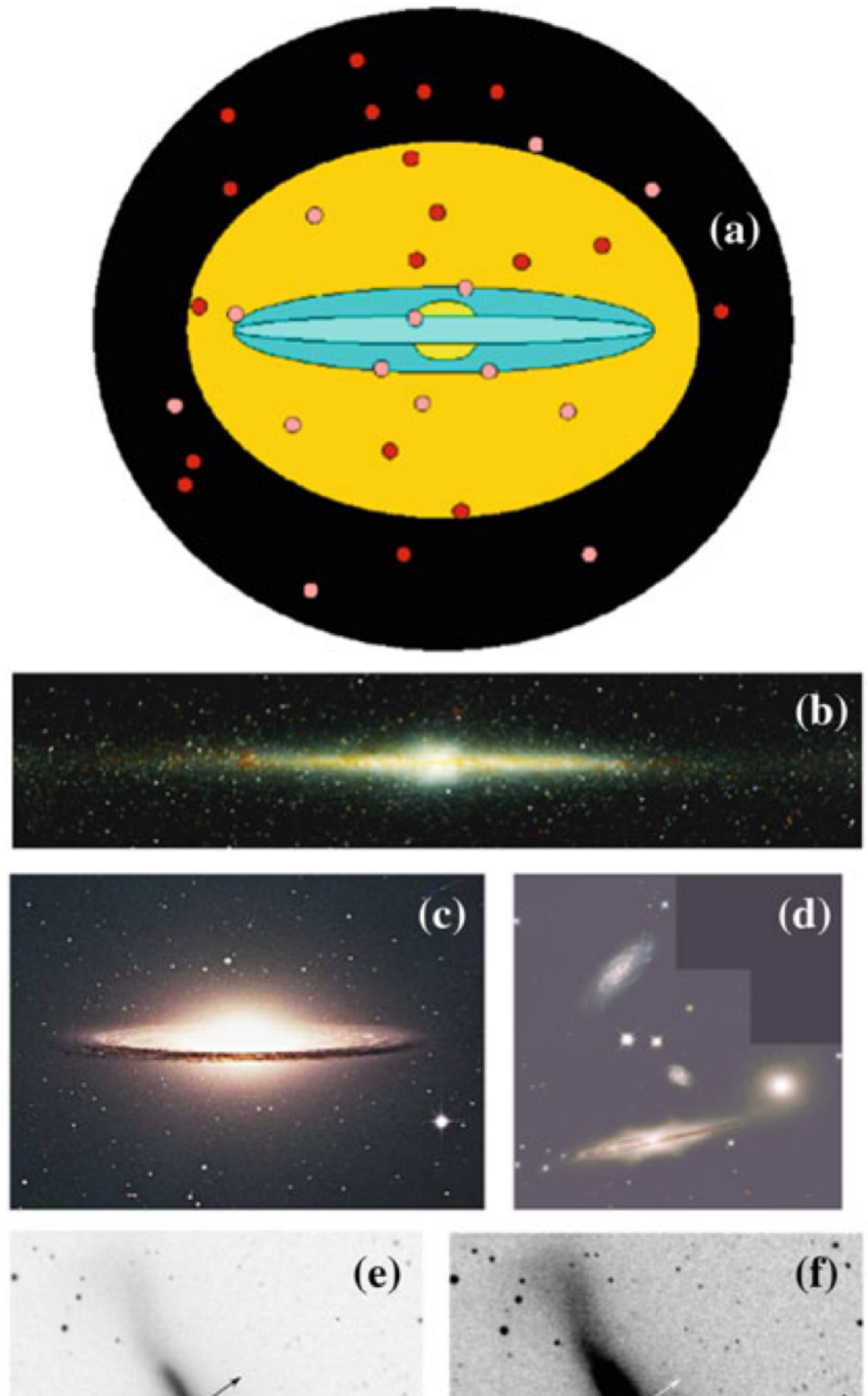


Fig. 1.19 **a** Sketch of Milky Way showing the stellar disk (*light blue*), thick disk (*dark blue*), stellar bulge (*yellow*), stellar halo (*mustard yellow*), dark halo (*black*) and globular cluster system (*filled circles*). The radius of the stellar disk is roughly 15 kpc. The baryon and dark halos extend to a radius of at least 100 kpc. **b** Infrared image of the Milky Way taken by the DIRBE instrument on board the Cosmic Background Explorer (COBE) Satellite (we acknowledge the NASA Goddard Space Flight Center and the COBE Science Working Group for this image). **c** M104, a normal disk galaxy with a large stellar bulge (from AAO). **d** Hubble Heritage image of the compact group Hickson 87; one galaxy has a peanut-shaped stellar bulge due to dynamical interaction with other group members. **e** Image of the SO galaxy NGC 4762 (Digital Sky Survey) shows its thin disk and stellar bulge. **f** A deeper image of NGC 4762 (DSS) shows

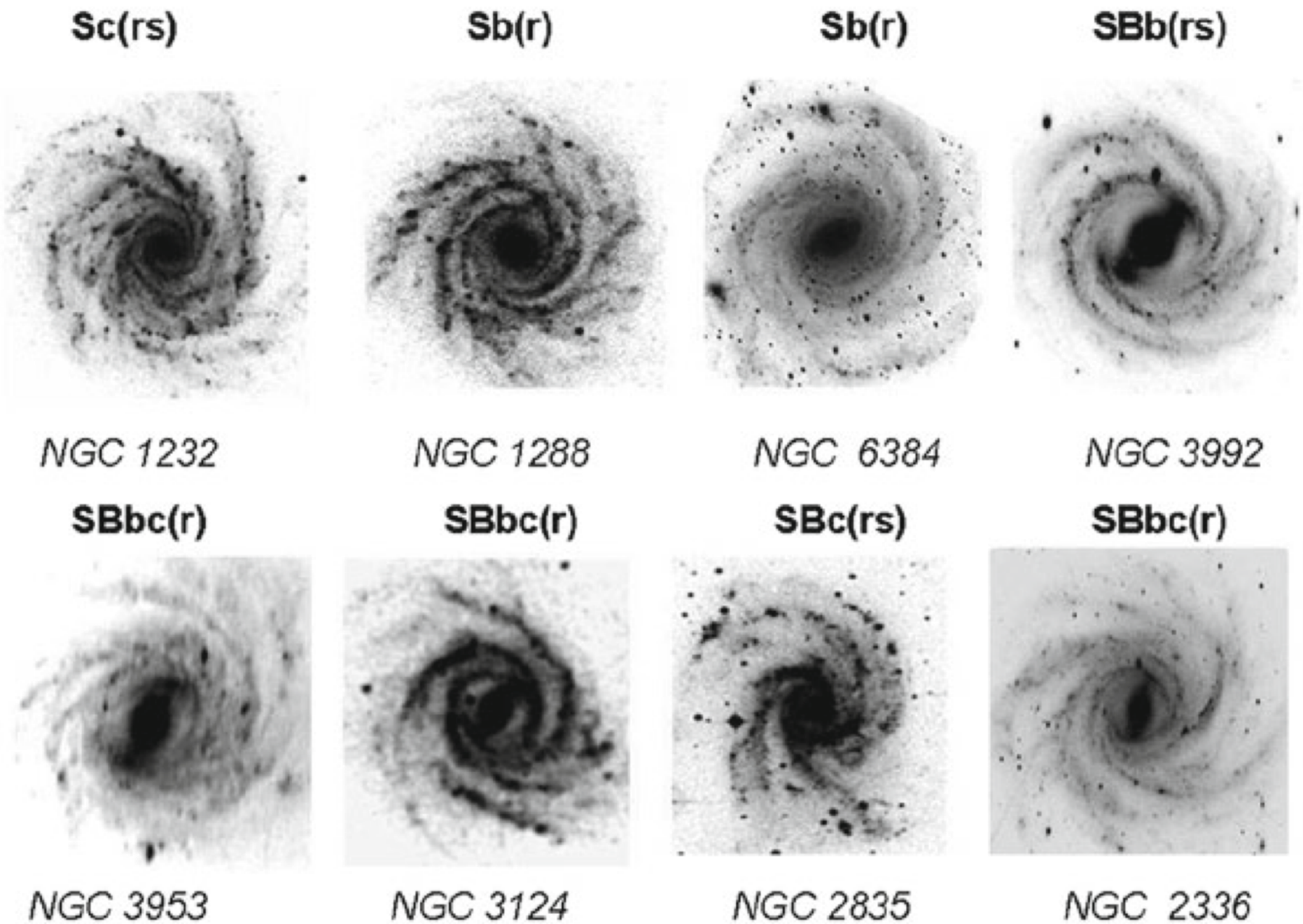


Fig. 1.18 A selection of galaxies that bear resemblance to the Milky Way (Efremov 2011); NGC 3992 is arguably the closest match to the Galaxy

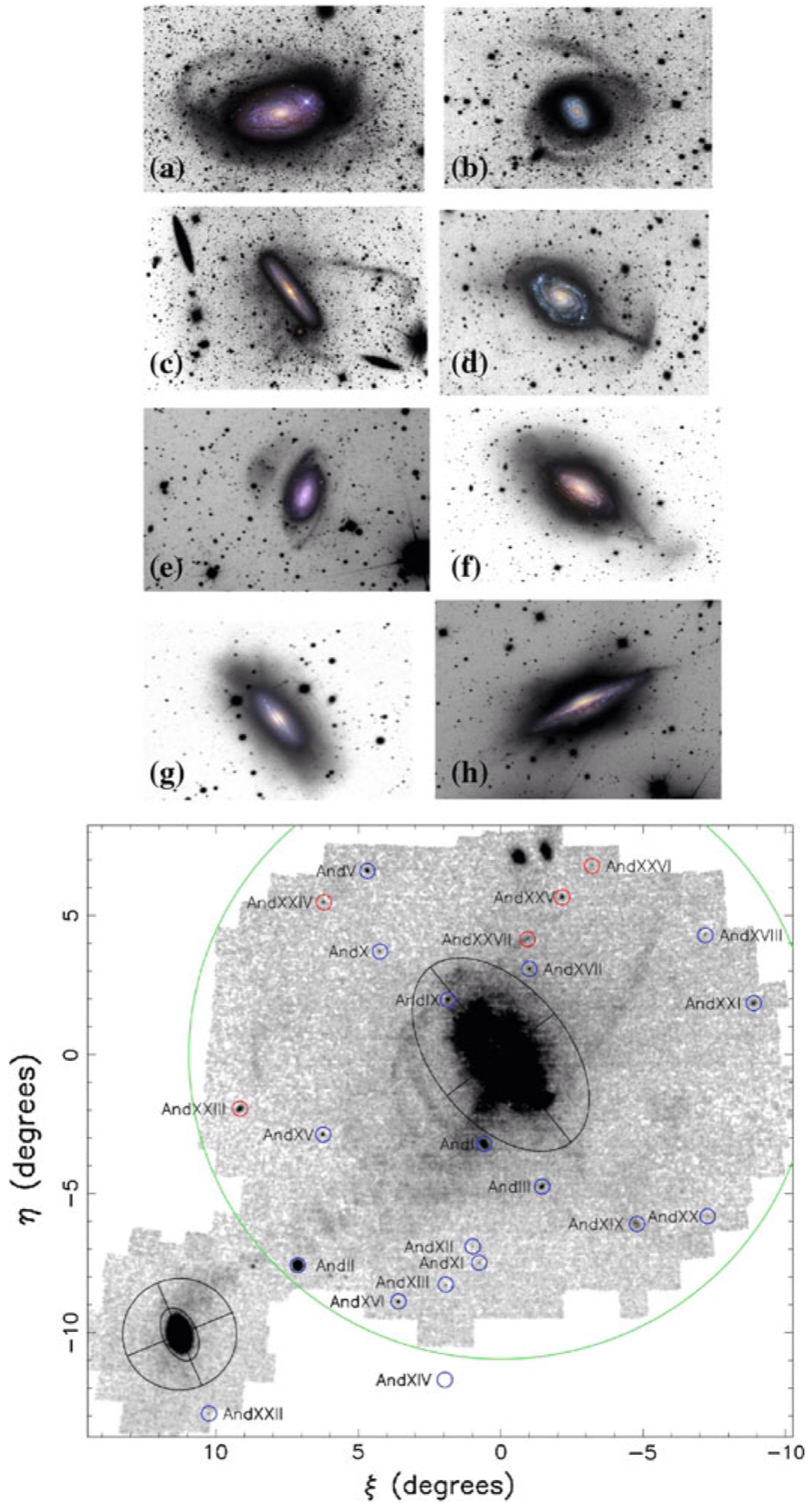


Fig. 1.20 *Upper* Examples of normal spirals with faint stellar streamers in the outer halo taken from a recent survey by Martinez-Delgado et al. 2010. *Lower* M31 and M33 system (PANDAs survey; Richardson et al. 2011); many associated dwarfs are also shown

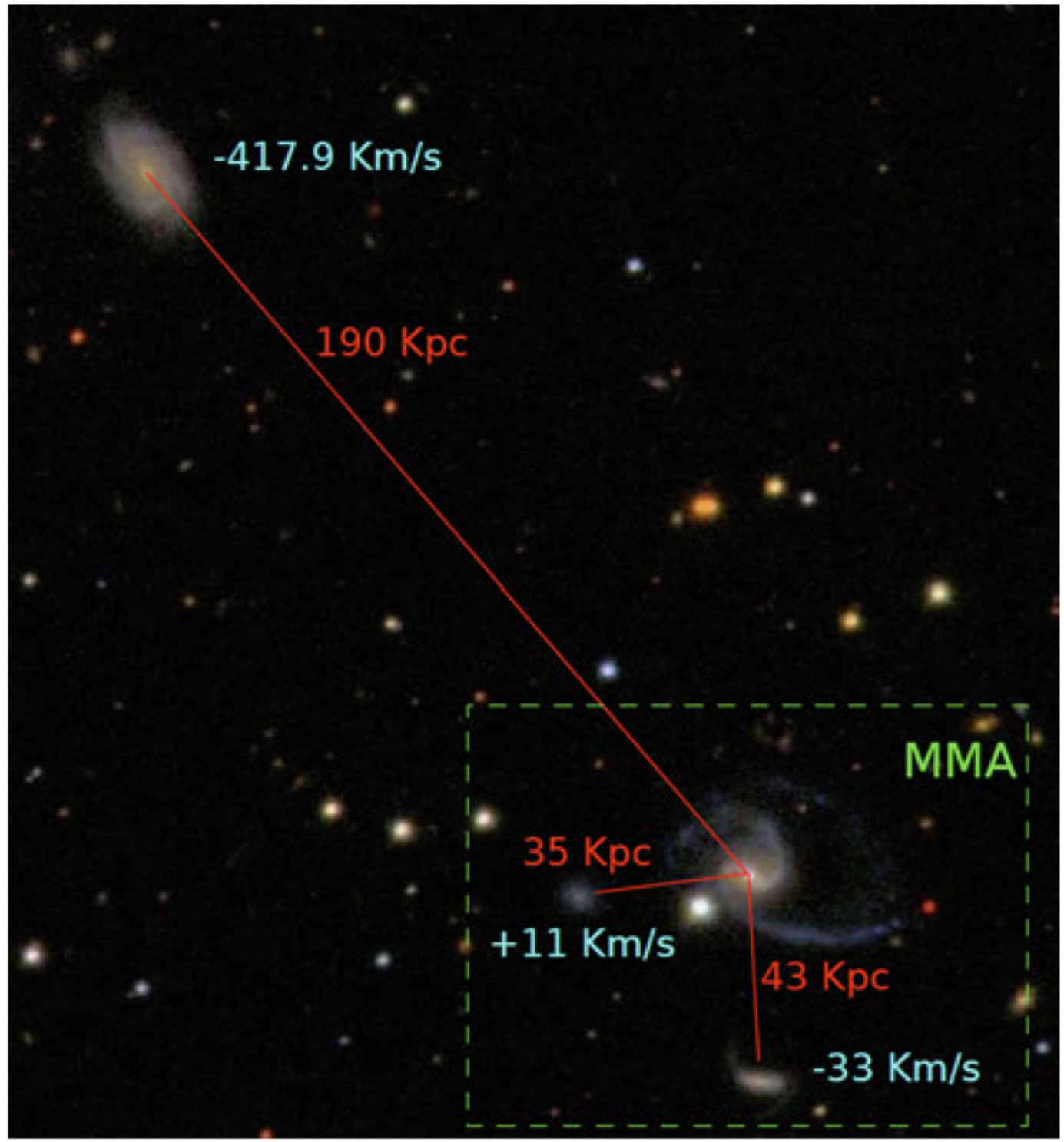


Fig. 1.9 An SDSS image of the best analogue of the Local Group to emerge from the GAMA galaxy survey (Robotham et al. 2012). The group marked MMA is remarkably similar to the Milky Way—Magellanic System where the three main galaxies are all forming stars at the present time

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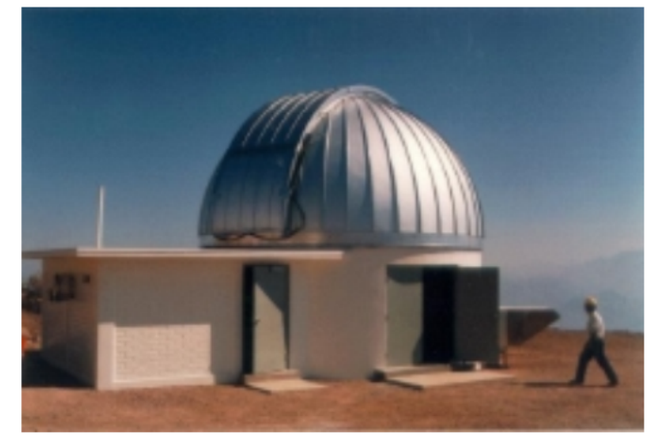
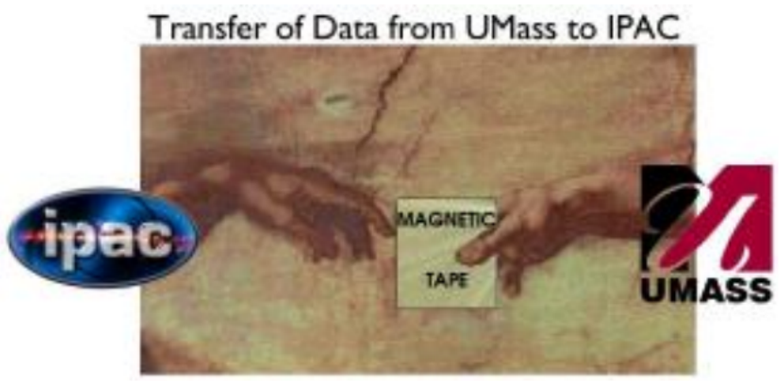
The two 2MASS Cassegrain-focus equatorial-mount reflector telescopes are identical in construction, and each has a primary mirror diameter of 1.3 meters. One telescope is located on Mt. Hopkins in Arizona (*left*), and the other on Cerro Tololo, along the Andes mountains in Chile (facility seen at *right*). Each telescope was equipped with a three-channel camera, each channel consisting of a 256-pixel × 256-pixel detector array, capable of observing the sky simultaneously at wavelength bands centered at wavelengths 1.25 micrometers, 1.65 micrometers, and 2.17 micrometers. (1 micrometer, abbreviated 1 μm, is one-millionth of a meter in length.) These bands are colloquially known by astronomers as "*J*", "*H*", and "*K-short*" (or "K_s"). The refrigerated dewar containing the camera is the brass canister hanging below the telescope. Computers operated each telescope's motion with respect to the sky; the motion of the much-smaller secondary mirror (seen near the top of the telescope); and the dark shutters of the three channels and digital "readout" of the three detector arrays.





An observing "schedule file" was sent down from UMass to the telescope operator at each of the two facilities for each night. Computers used these files to automatically control the telescopes throughout the night. The operator at each facility monitored the telescope operations and the acquiring of data. Each morning the digital data were written to magnetic tape storage, and the tapes were sent to IPAC. Each night's data were put through an automated "*pipeline*" of computer programs, which convert the raw images made by the telescopes into final *processed* images and catalogs containing star and galaxy brightnesses and positions. These data were assessed for quality, through some human intervention. If these data passed muster, they were included as part of the enormous 2MASS *database*, from which a *digital image atlas of the sky* and *catalogs* of point sources and extended sources, freely available to astronomers and the public at large, were assembled and distributed.

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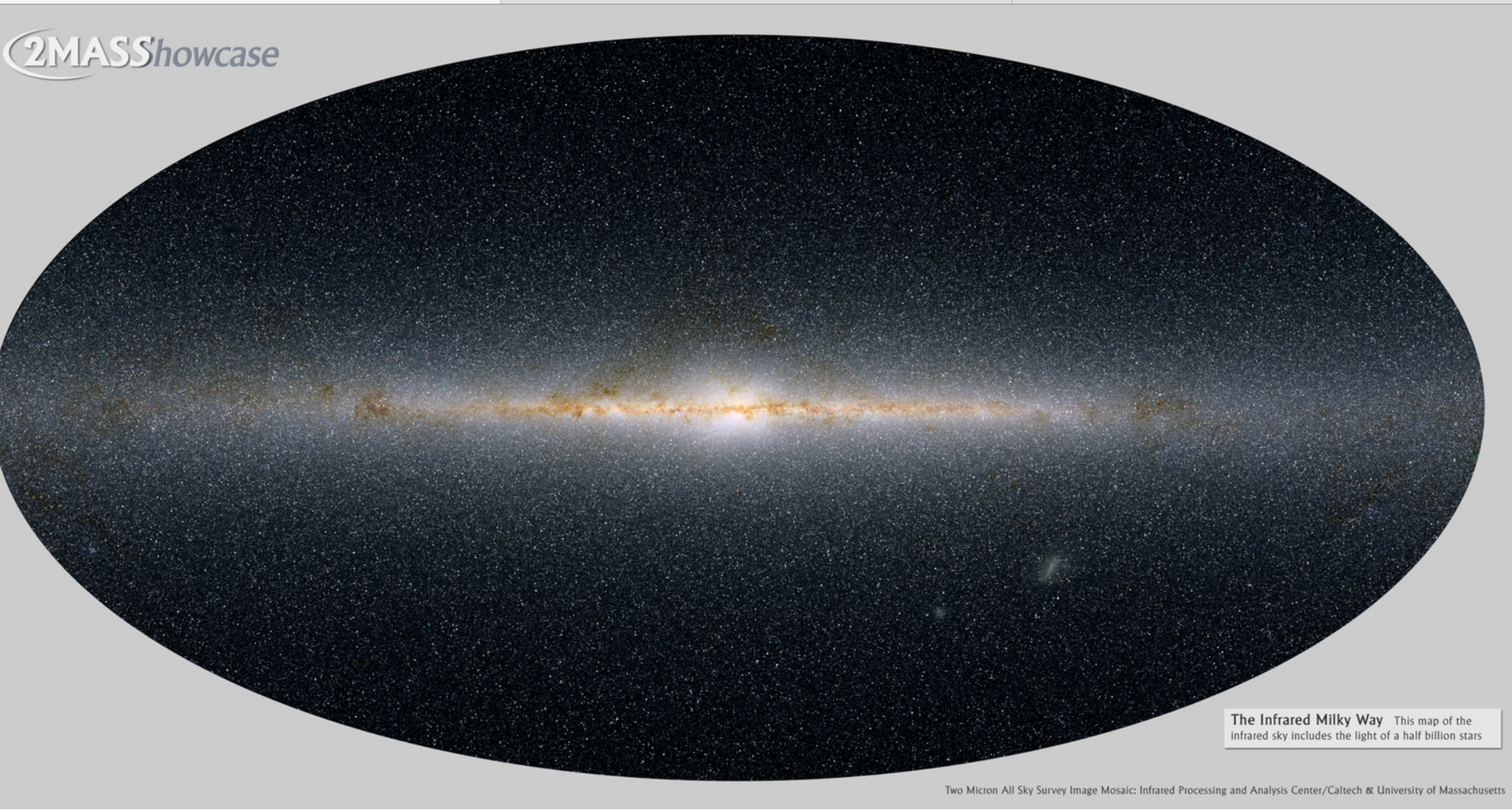
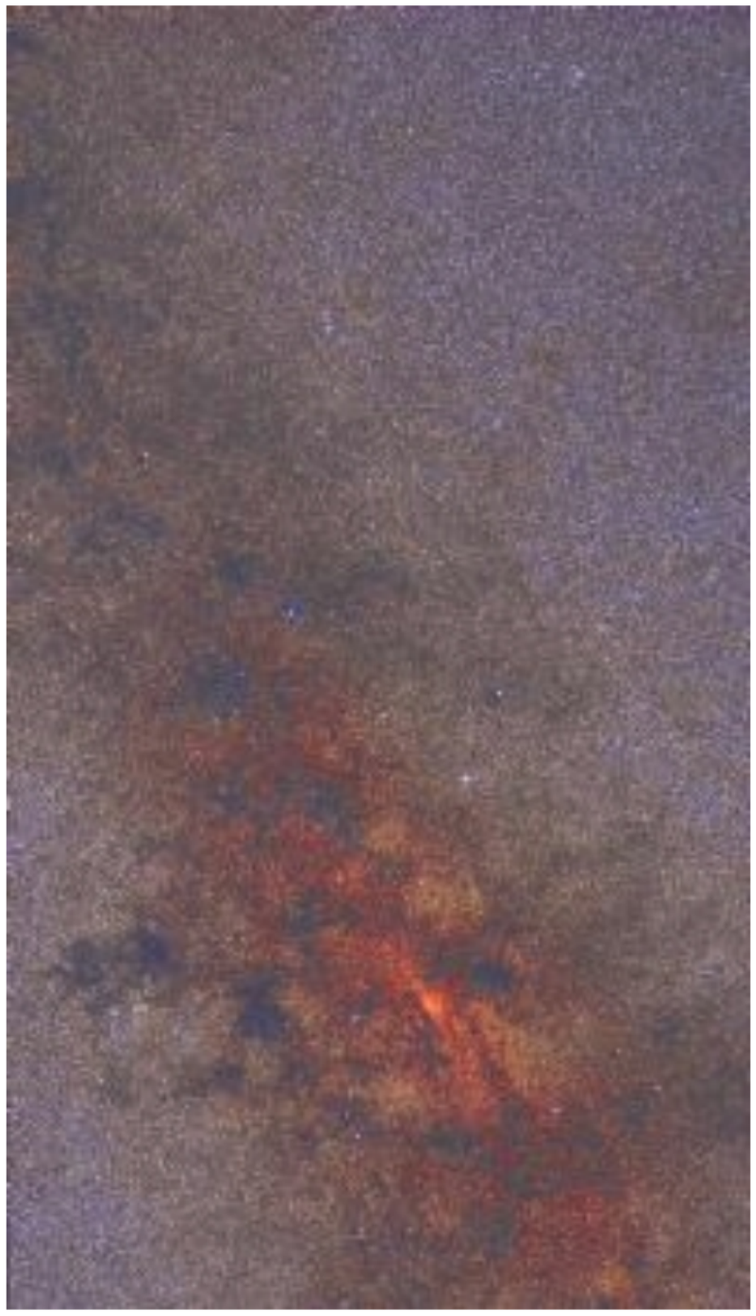
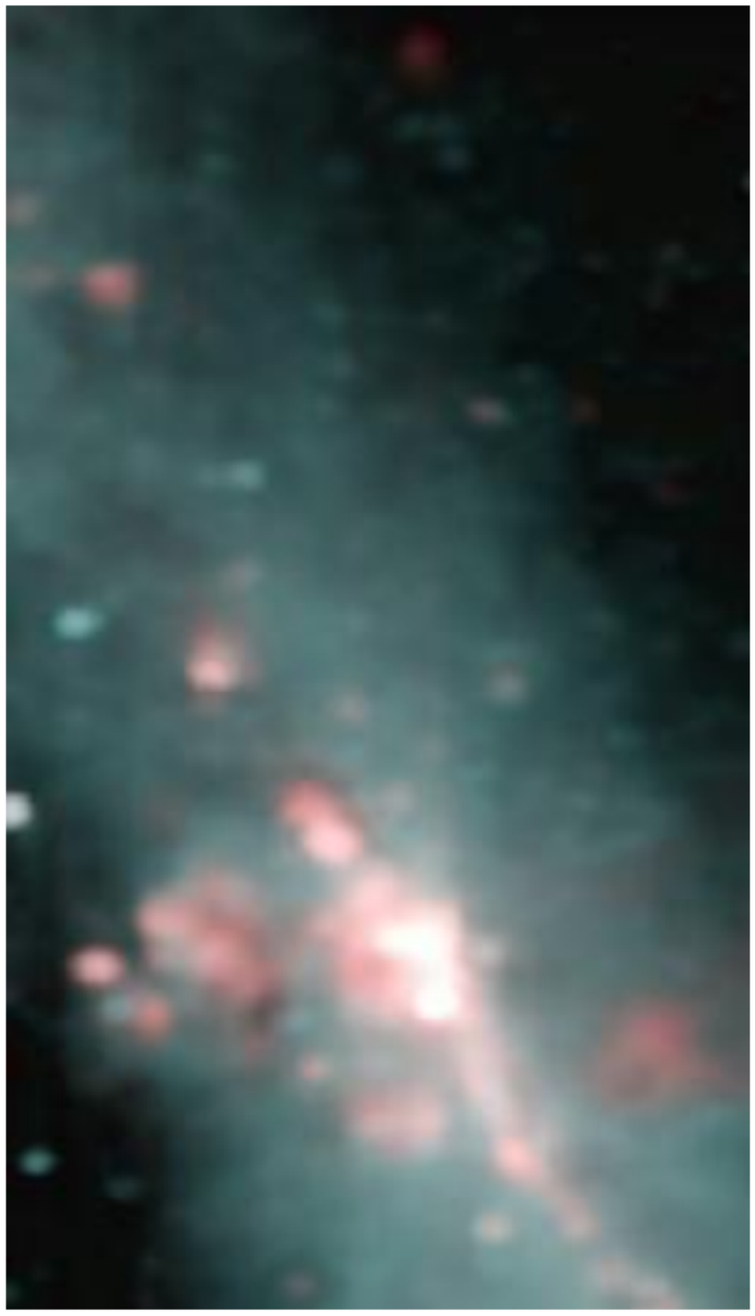


The Infrared Milky Way This map of the infrared sky includes the light of a half billion stars

Two Micron All Sky Survey Image Mosaic: Infrared Processing and Analysis Center/Caltech & University of Massachusetts

POSS	2MASS	IRAS
		
Optical	Near-Infrared	Far-Infrared

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Sloan Digital Sky Survey | milky way - Google Search | 1962ApJ...136..748E

Sloan Digital Sky Survey

Mapping the Universe

The Sloan Digital Sky Survey

The Sloan Digital Sky Survey (SDSS) is one of the most ambitious and influential surveys in the history of astronomy. Over eight years of operations (SDSS-I, 2000-2005; SDSS-II, 2005-2008), it obtained deep, multi-color images covering more than a quarter of the sky and created 3-dimensional maps containing more than 930,000 galaxies and more than 120,000 quasars.

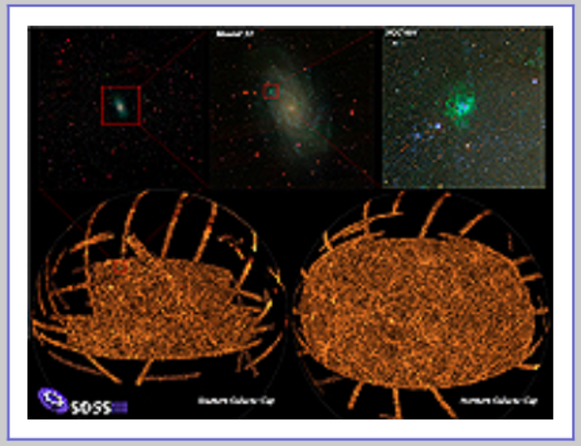
SDSS data have been released to the scientific community and the general public in annual increments, with the final public data release from SDSS-II occurring in October 2008. That release, Data Release 7, is available through this website.

Meanwhile, SDSS is continuing with the Third Sloan Digital Sky Survey (SDSS-III), a program of four new surveys using SDSS facilities. SDSS-III began observations in July 2008 and released Data Release 8 in January 2011, Data Release 9 in August 2012, and Data Release 10 in July 2013. SDSS-III will continue operating and releasing data through 2014.

Data Release 10 contains the first release of APOGEE infrared Galactic spectroscopy as well as cumulative updates to the BOSS optical extragalactic spectroscopy archive.

Data Release 9 contains the first release of BOSS spectroscopy to the public as well as

Images of the SDSS
(click for more information)



The Final Survey

The Whirlpool Galaxy (M51)

Sloan Digital Sky Survey

- Új-Mexikó, 2,5m-es teleszkóp Apache Pointban
- Öt színben képalkotás több mint 100 millió égitestről
- Kb. 700 ezer spektrum galaxisokról, kvazárokról és csillagokról
- Fontos magyar részvevők (Szalay Sándor, Csabai István és tanítványaik)

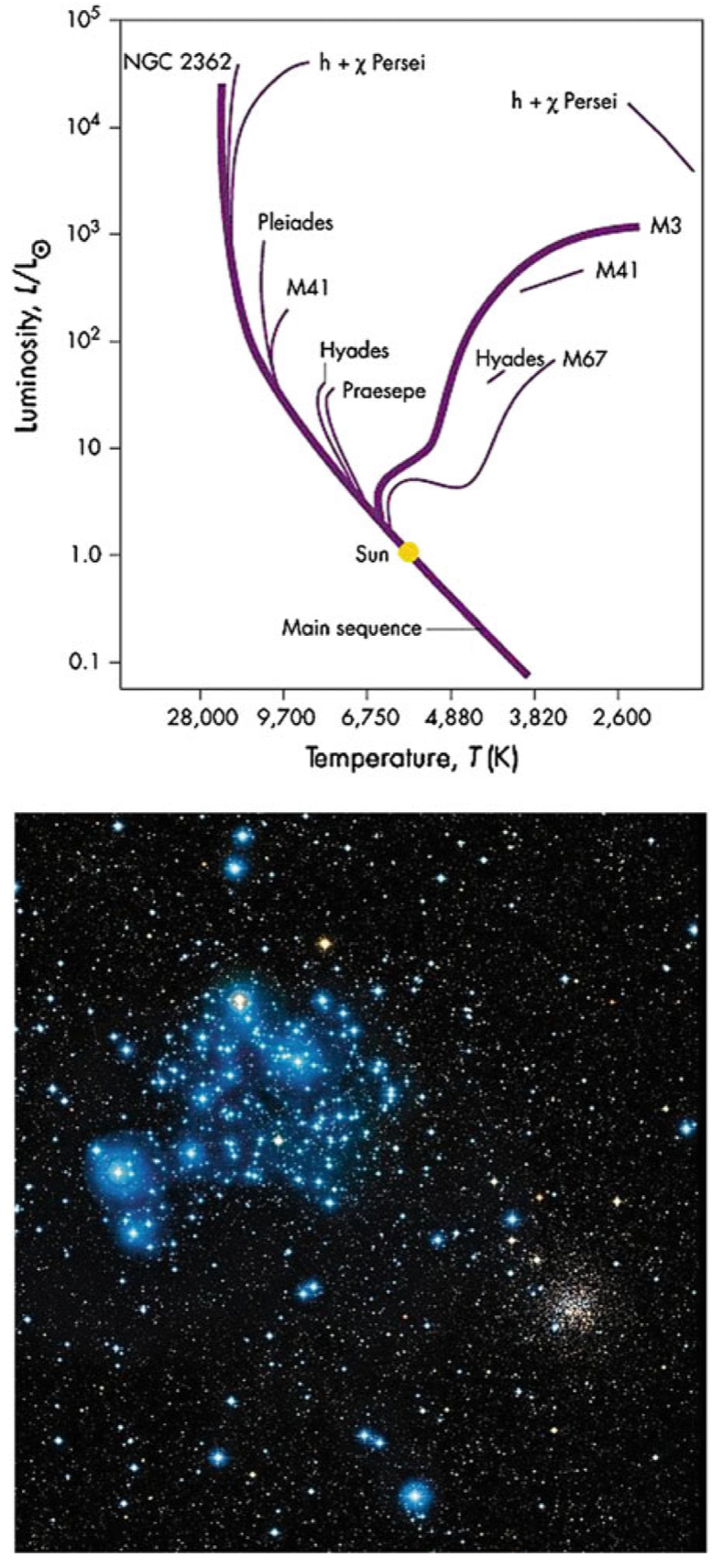


Fig. 1.21 *Upper* Hertzsprung-Russell diagram for stellar clusters over a wide range in age. *Lower* CFHT imaging of two open clusters M35 (NE) and NGC 2158 (SW) in Gemini. The young cluster M35 (NGC 2168) has a mass of roughly 2500 $M_\odot$ and is 850 pc from the Sun. The massive cluster NGC 2158 is four times further away and about 1 Gyr old

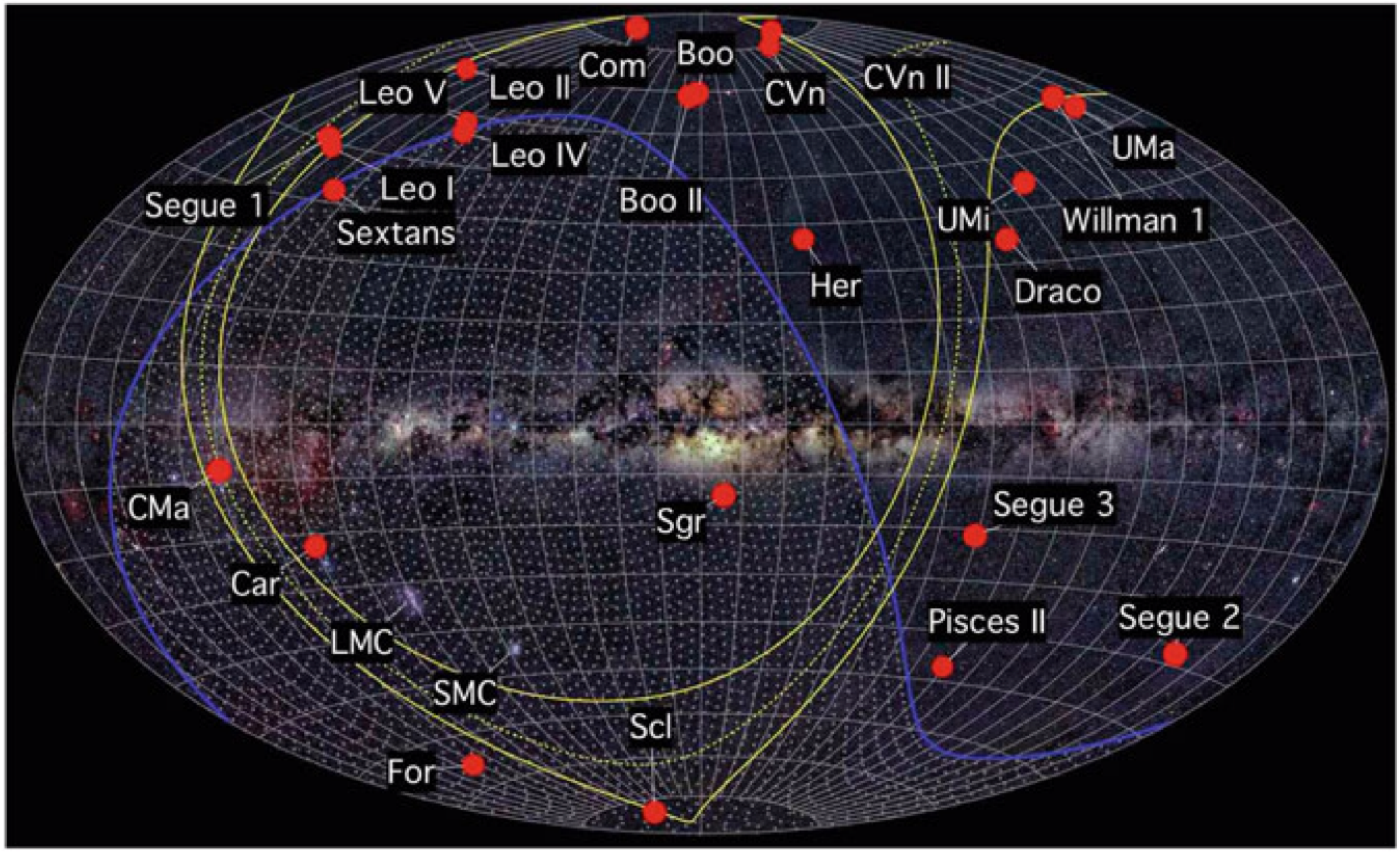


Fig. 1.22 *Upper* Aitoff projection of a satellite in orbit about the Milky Way as it would appear after 8 Gyr. While stars from the disrupted satellite appear to be dispersed over a very wide region of sky, it will be possible to deduce the parameters of the original event using special techniques (see text) (we acknowledge A. Helmi and S. White for this image). *Lower* Aitoff projection of most of the known dwarfs in orbit about the Galaxy. For decades, astronomers have argued that these are broadly confined to a plane rather than spherically distributed, an effect that is also seen in the PANDAs survey for the dwarfs in M31 (we acknowledge A. Frebel for this image)

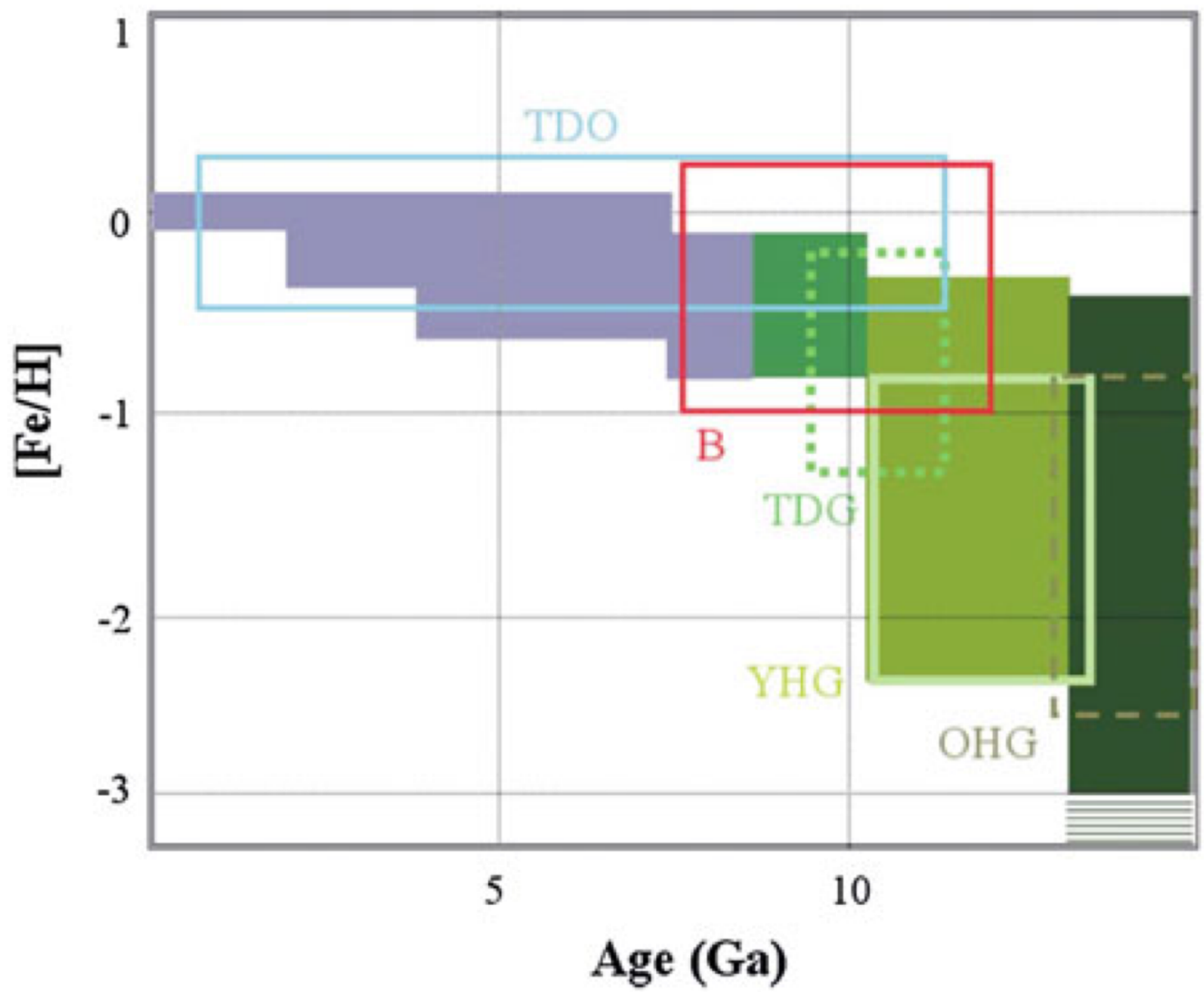


Fig. 1.16 The age-metallicity relation of the Galaxy for the different components (see text): *TDS* thin disk stars; *TDO* thin disk open clusters; *ThDS* thick disk stars; *ThDG* thick disk globulars; *B* bulge; *YHG* young halo globulars; *OHG* old halo globulars

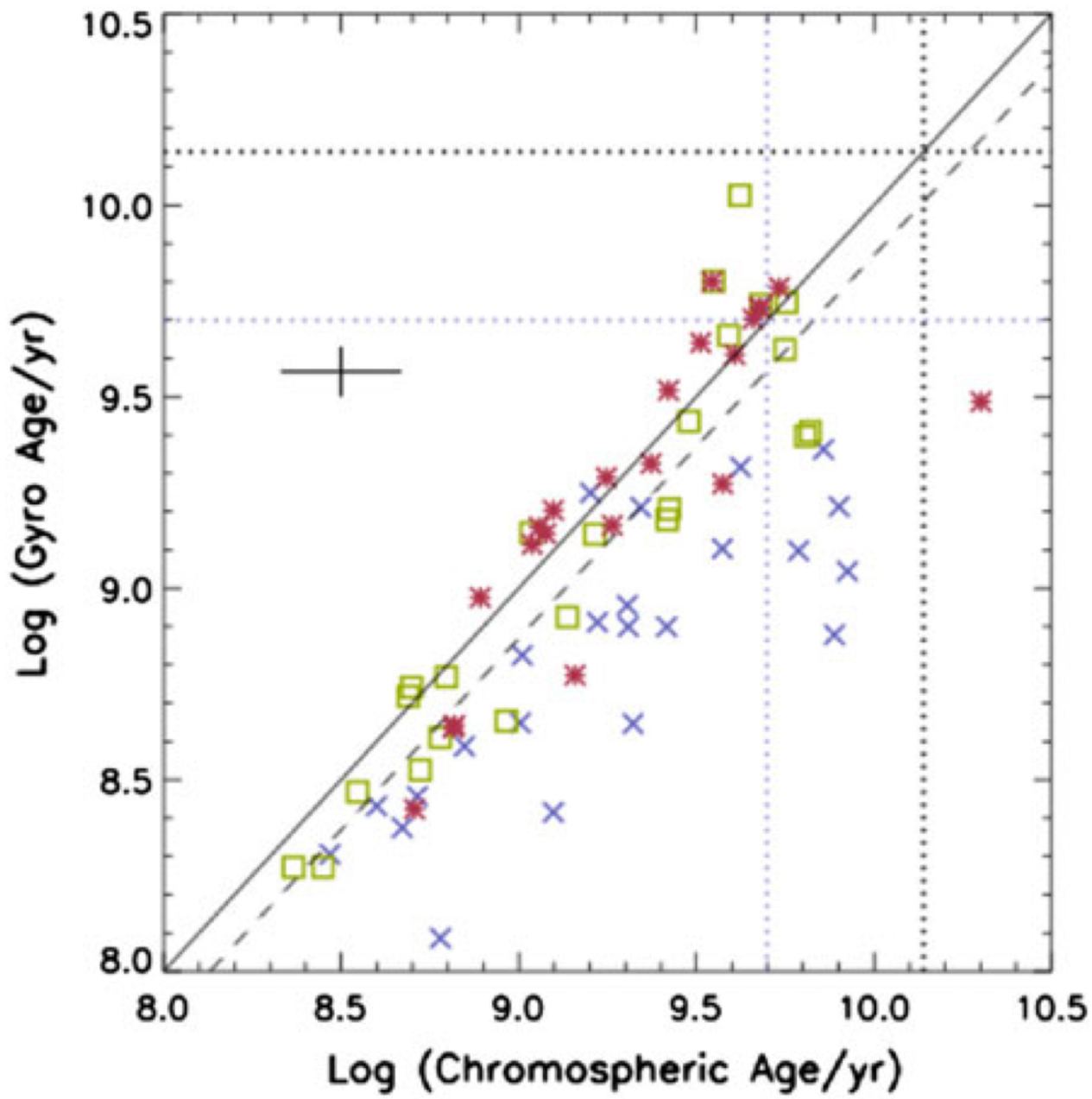


Fig. 1.17 Gyro ages versus chromospheric ages for a sample of well-studied stars (Barnes 2007). The *blue crosses* are for stars bluer than $B - V = 0.6$ and should be disregarded

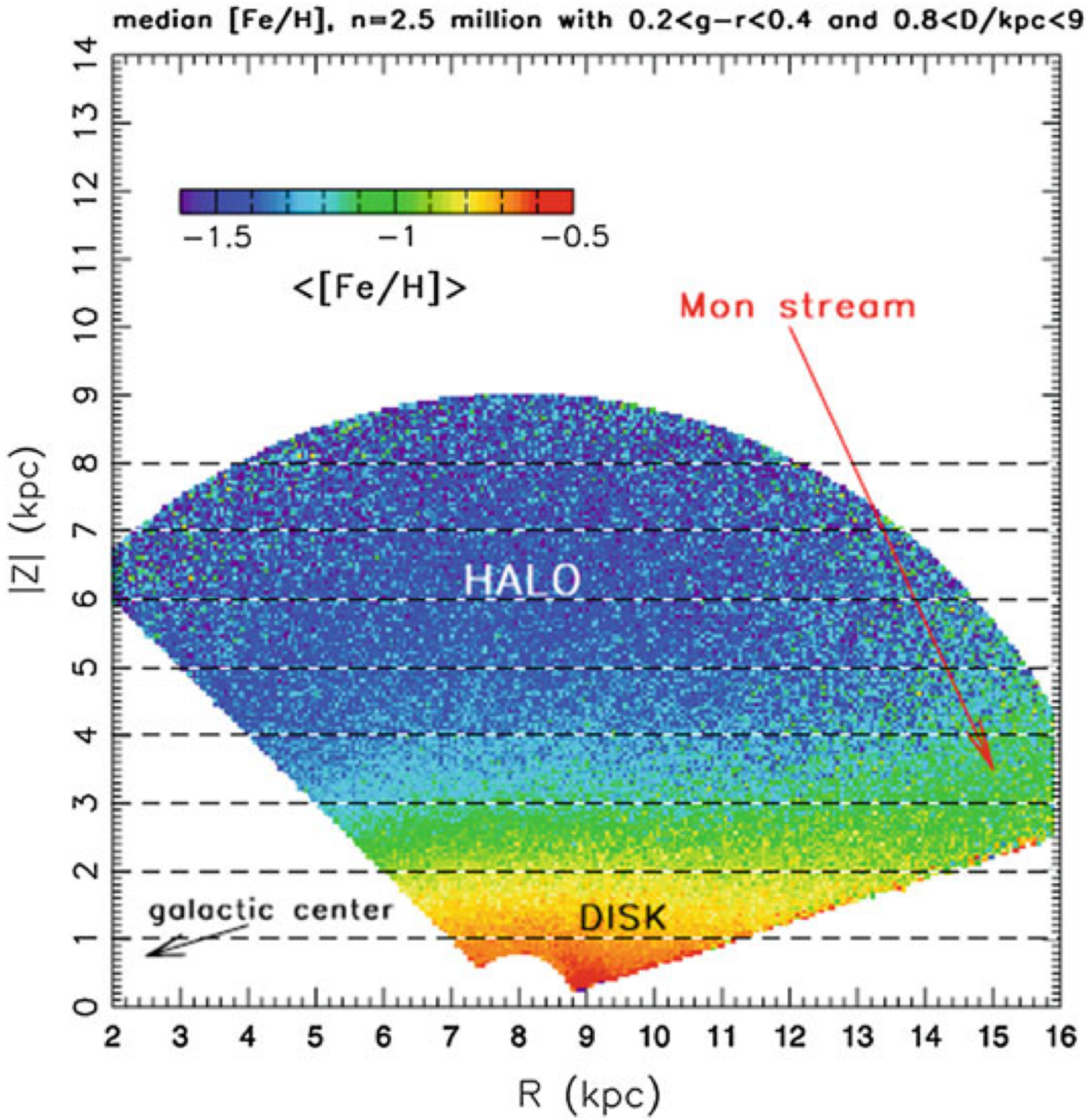


Fig. 1.32 Distribution of metallicity in the Galaxy from SDSS photometry (from Ivezic et al. 2008)

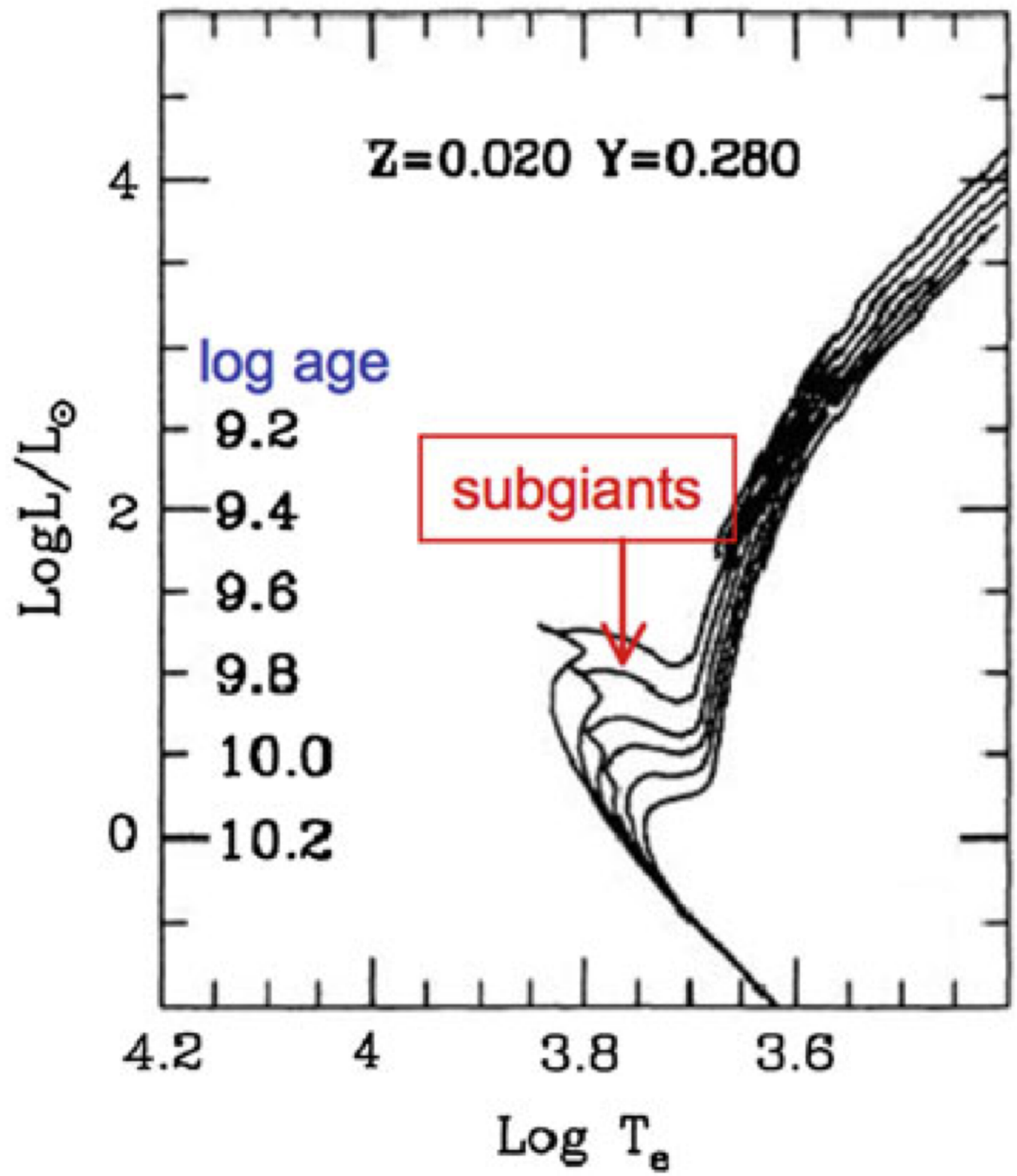
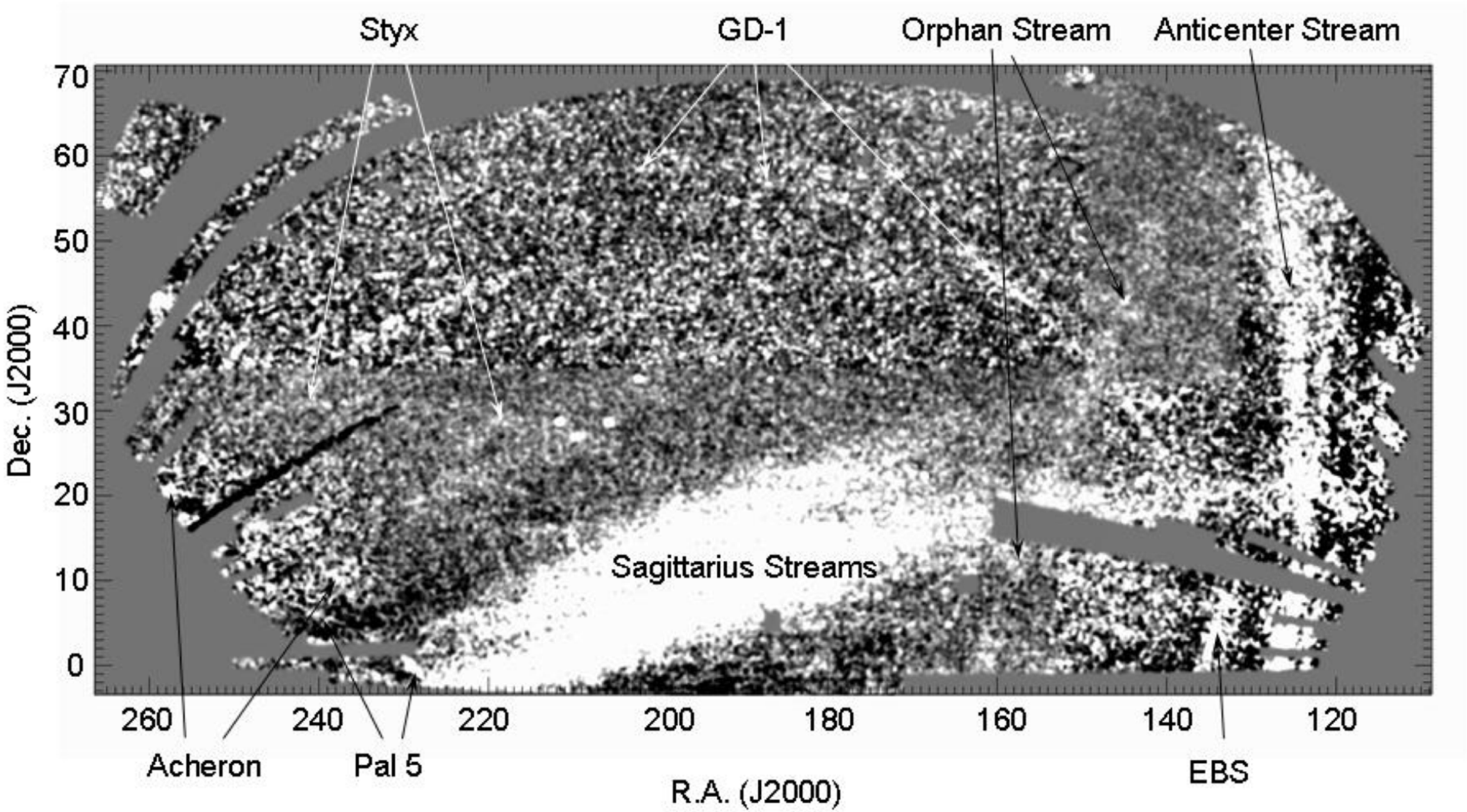


Fig. 1.39 Isochrones from Bertelli et al. (1994) for stars of near-solar abundance and a range of ages from 1.6 to 16 Gyr. The isochrones can overlap near the turnoff are well separated in the subgiant region

Table 1 Known Distant Galactic Streams

Designation	Progenitor	Selected References
Sagittarius	Sagittarius Dwarf Galaxy	Ibata et al. 1994, Mateo et al. 1996, Alard 1996, Toten & Irwin, 1998, Ibata et al. 2001, Martinez-Delgado et al. 2004, Majewski et al. 2004, Vivas et al. 2005, Belokurov et al. 2006b, Fellhauer et al. 2006, Bellazzini et al. 2006, Chou, M-Y et al. 2007, Law et al. 2009
Virgo Stellar Stream	NGC 2419?	Vivas et al. 2001, Duffau et al, 2006, Newberg et al. 2007
Palomar 5	Palomar 5	Odenkirchen et al. 2001, 2003, 2009, Rockosi et al. 2002, Grillmair & Dionatos 2006b
Monoceros Ring	Unknown (dwarf galaxy?)	Newberg et al. 2002, Yanny et al. 2003, Ibata et al. 2003, Rocha-Pinto et al. 2003, Penarrubia et al. 2005
NGC 5466	NGC 5466	Belokurov et al. 2006a, Grillmair & Johnson 2006, Fellhauer 2007
Orphan Stream	Unknown (dwarf galaxy?)	Grillmair 2006a, Belokurov 2007 Fellhauer et al. 2007, Sales et al. 2008, Newberg et al. 2010
GD-1	Unknown (globular cluster?)	Grillmair & Dionatos 2006b, Willet et al. 2009, Koposov, Rix, & Hogg 2009
AntiCenter Stream	Unknown (dwarf galaxy?)	Grillmair 2006b, Grillmair, Carlin, & Majewski 2008
EBS	Unknown (dwarf galaxy?)	Grillmair 2006, Grillmair, Carlin, & Majewski 2008
Acheron	Unknown (globular cluster?)	Grillmair 2009
Cocytos	Unknown (globular cluster?)	Grillmair 2009
Lethe	Unknown (globular cluster?)	Grillmair 2009
Styx	Bootes III dwarf?	Grillmair 2009
Cetus Polar Stream	NGC 5824?	Newberg, Yanny, & Willett 2009


Styx
GD-1
Orphan Stream
Anticenter Stream
Sagittarius Streams
Acheron
Pal 5
EBS
Dec. (J2000)
R.A. (J2000)
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40
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10
0
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240
220
200
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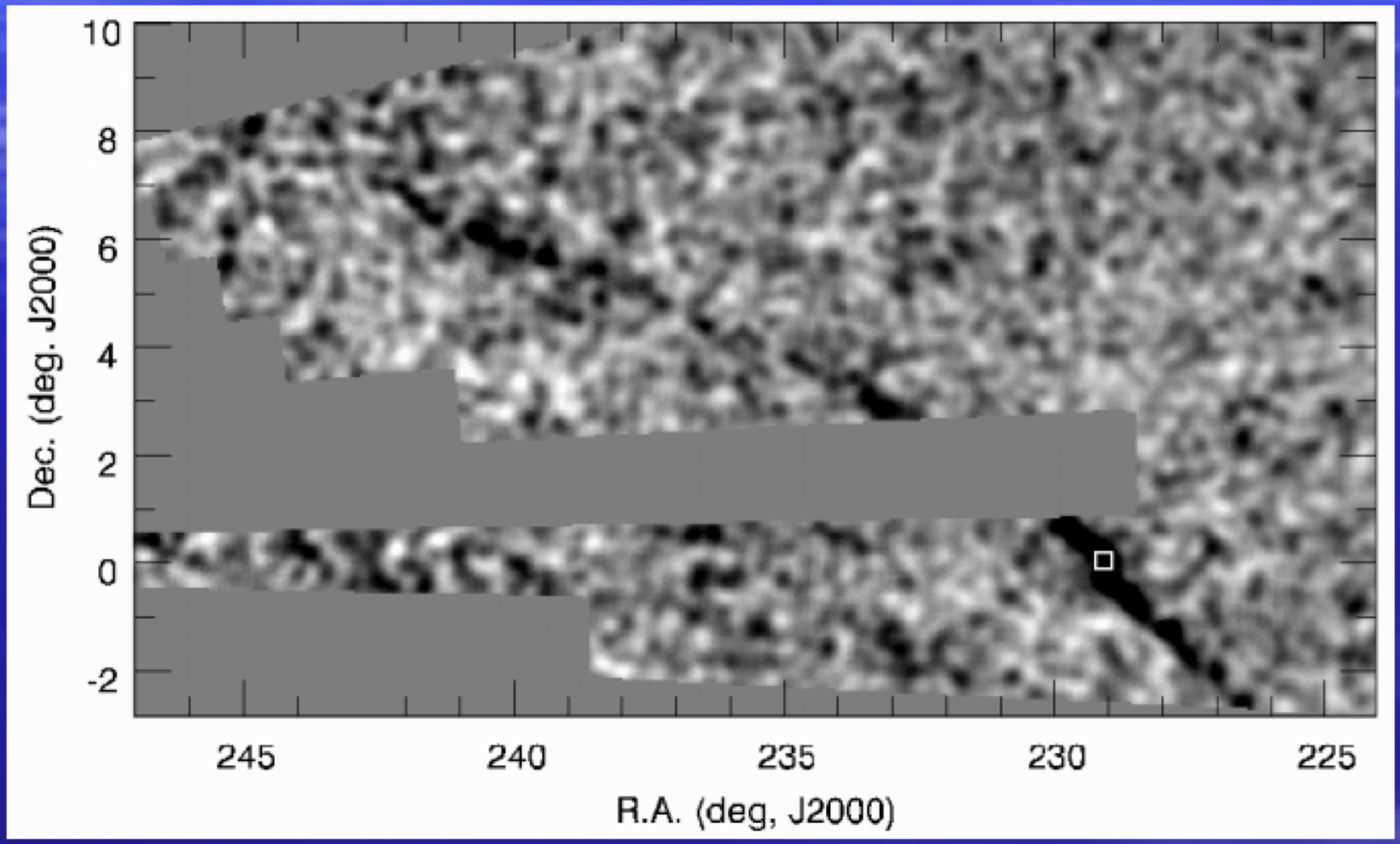
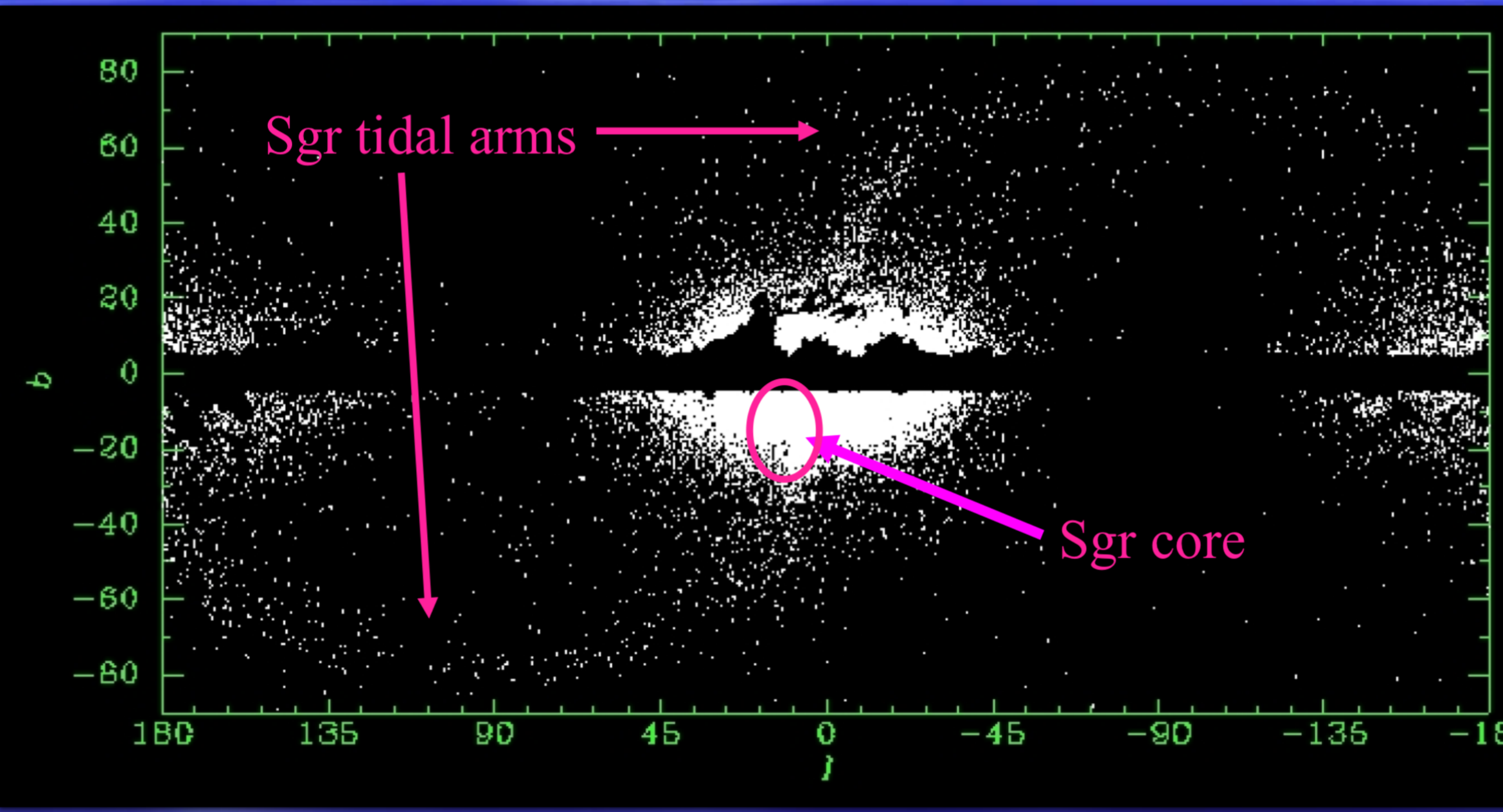
160
140
120



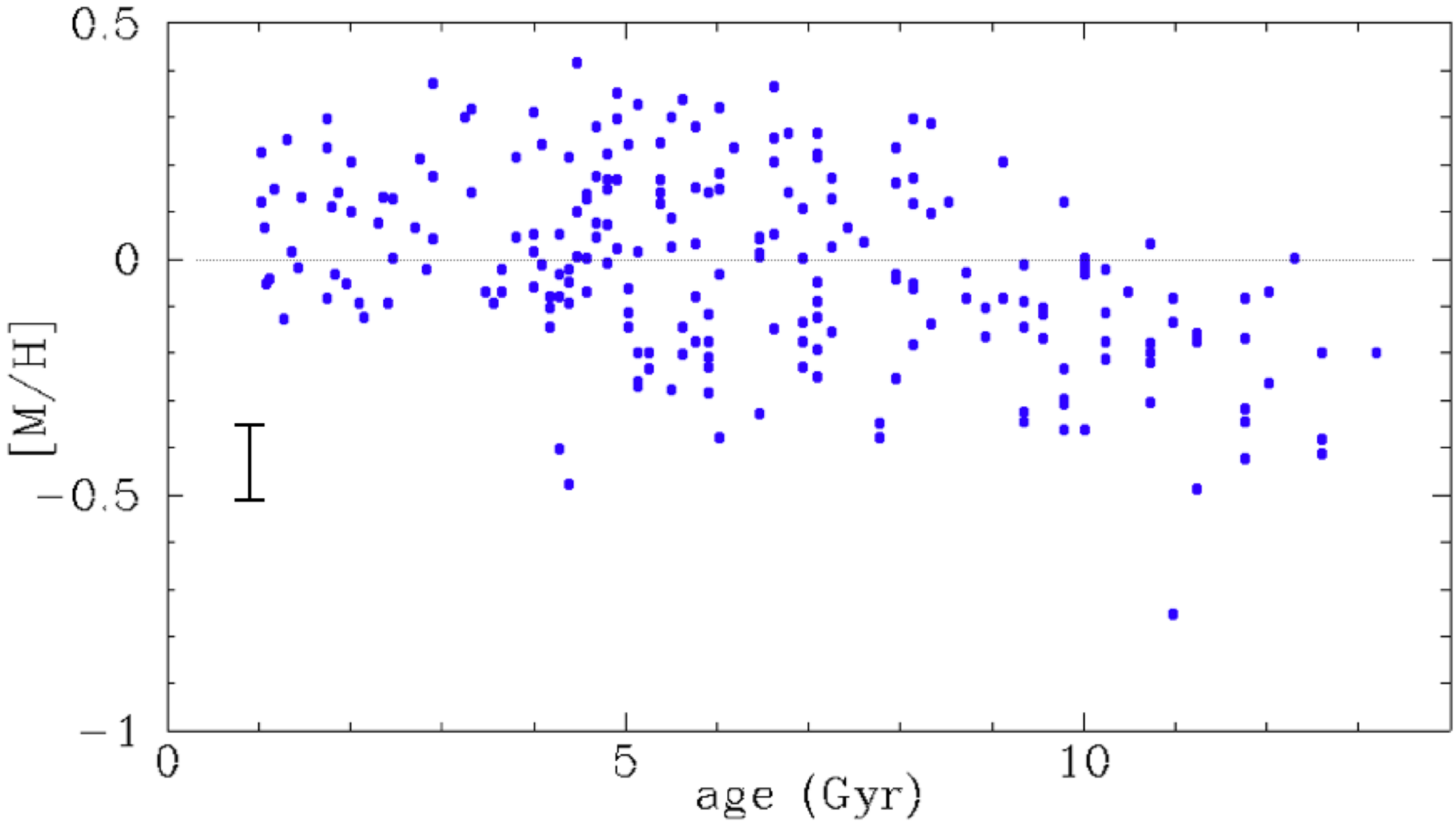
Grillmair & Dionatos 2006

Sagittarius dSph cont' d

All-sky view of 2MASS M giant star distribution



Majewski et al. 2003



1: The age-metallicity relation for subgiants near the sun (Wylie de Boer & Freeman
shed). The estimated random errors in [M/H] are shown by the error bar. The random
n age are believed to be about 30%.

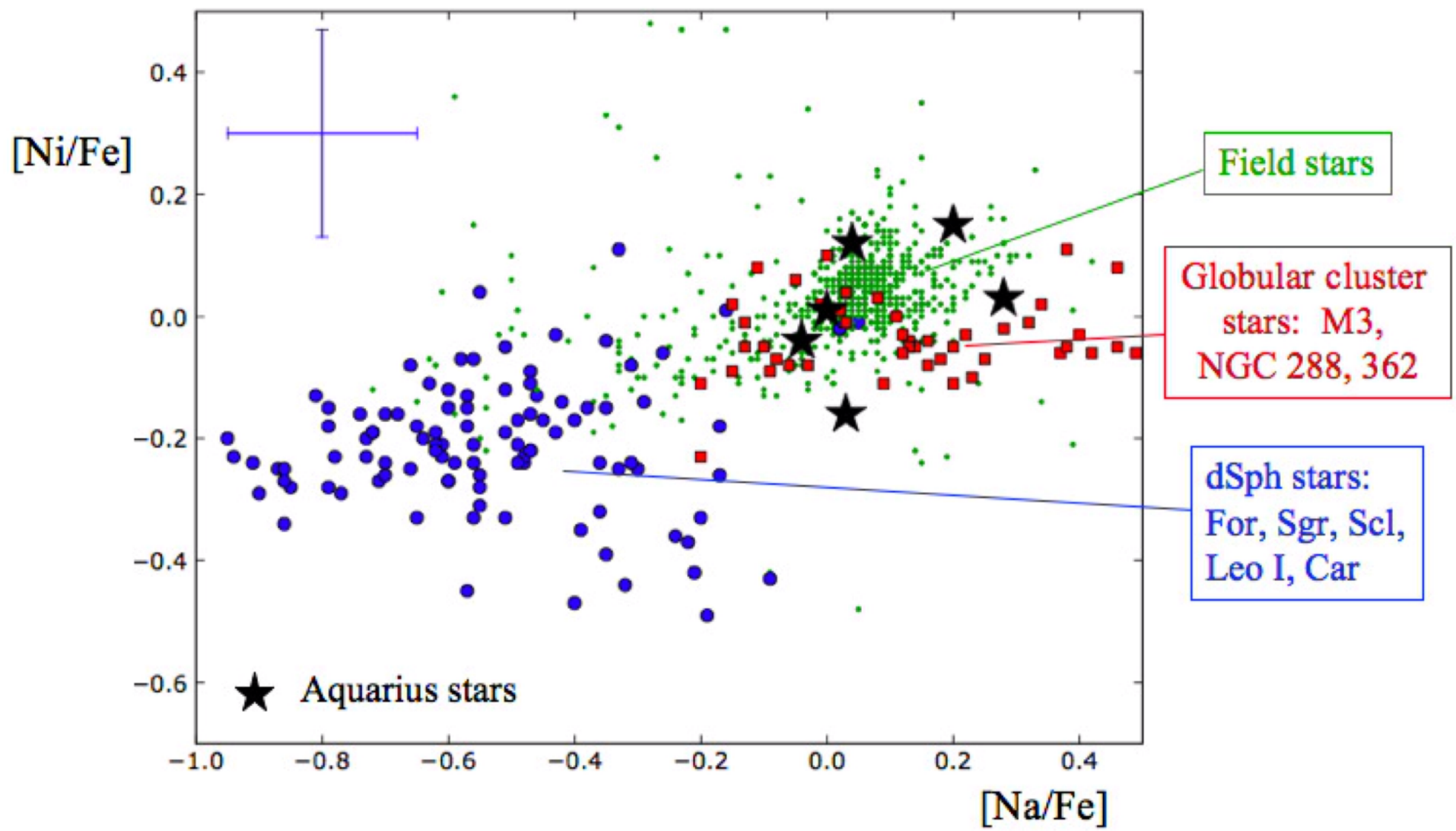
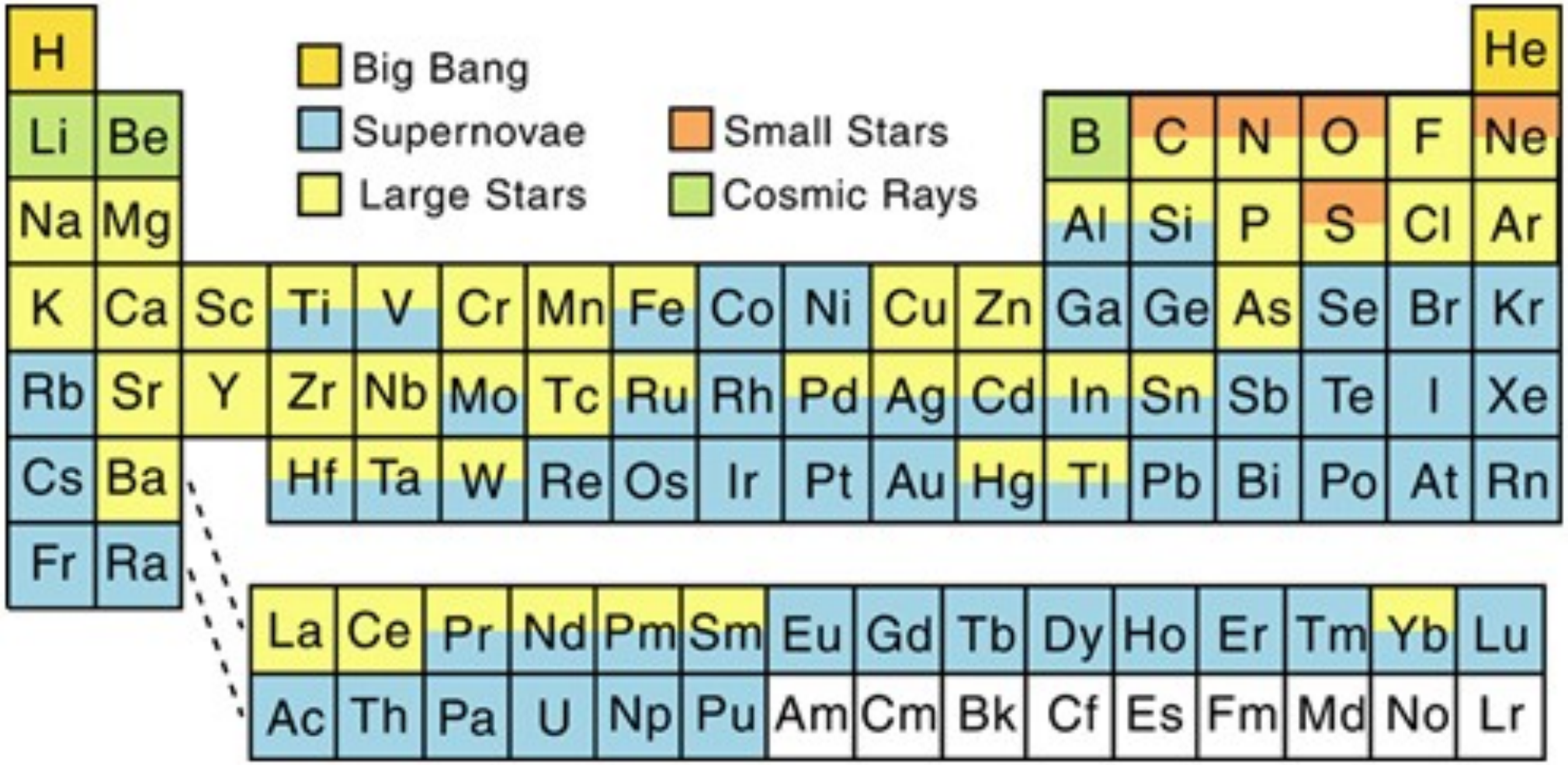


Figure 4: The Na-Ni distribution for globular cluster stars, dwarf spheroidal galaxy stars, field halo stars and stars of the Aquarius stream (black star symbols) (Wylie de Boer et al. 2012). The stars of the Aquarius stream are in the same part of the distribution as the globular cluster stars.

galaxies. The distribution of their stars in the [X/Fe]–[Fe/H] plane is well defined for an individual galaxy but differs in structure from galaxy to galaxy, depending on their star formation history


Big Bang
Supernovae
Large Stars
Small Stars
Cosmic Rays
H He
Li Be B C N O F Ne
Na Mg Al Si P S Cl Ar
K Ca Sc Ti V Cr Mn Fe Co Ni Cu Zn Ga Ge As Se Br Kr
Rb Sr Y Zr Nb Mo Tc Ru Rh Pd Ag Cd In Sn Sb Te I Xe
Cs Ba Hf Ta W Re Os Ir Pt Au Hg Tl Pb Bi Po At Rn
Fr Ra
La Ce Pr Nd Pm Sm Eu Gd Tb Dy Ho Er Tm Yb Lu
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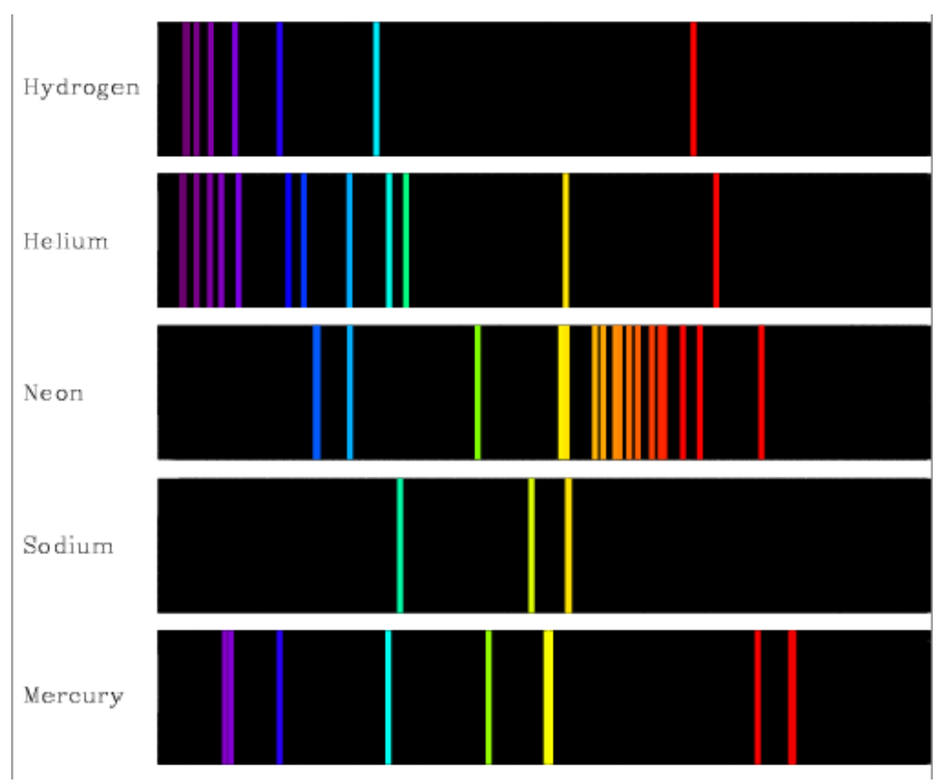
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Sloan Digital Sky Survey WEAVE 1962ApJ...136..748E

MOS with multiplex > 300

Project	Hemisphere	Tel. diam. (m)	FOV (sq deg)	Fibres / slits	Spectroscopic resolution	First light
AAT / AAOmega	S	3.9	3.14	392 fibres	1300 - 8000	2006
AAT / Hermes	S	3.9	3.14	400 fibres	28000	2012
CFHT / GYES	N	3.6	0.64	500 fibres	20000	Proposed
CTIO Blanco / DESpec	S	4.0	3.8	4000 fibres	2000	2012
GTC / GO-IRS	N	10.4	0.05	1000 fibres	5000 - 20000	2013
LAMOST	N	4.0	19.63	4000 fibres	1000, 10000	2009
Mayall / BigBOSS	N	3.8	7.07	5000 fibres	3000 - 4000	2016
MMT / HECTOSPEC	N	6.5	0.79	300 fibres	1000	2003
NTT or VISTA / 4MOST	S	4	7.1	3000 fibres	3000 - 5000	2017 - 2018
SDSS-III / BOSS	N	2.5	7.07	1000 fibres	1600 - 2700	2009
Subaru / PFS	N	8.2	1.78	2400 fibres	2000 - 5000	-
VLT / MOONS	S	8.2	0.14	500 fibres	3000 - 20000	2017
WHT / WEAVE	N	4.2	3.14	1000 fibres	5000, 20000	2016

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GALAH: Projects milky way - Google Search 1962ApJ...136..748E

GALAH GALactic Archaeology with HERMES

With GALAH, we will chemically tag stars into coeval groups, identifying individual members of star clusters which have long since dispersed. Using the stellar relics of ancient star formation and accretion events, we can reconstruct the Galactic accretion history, and dynamical and chemical evolution.

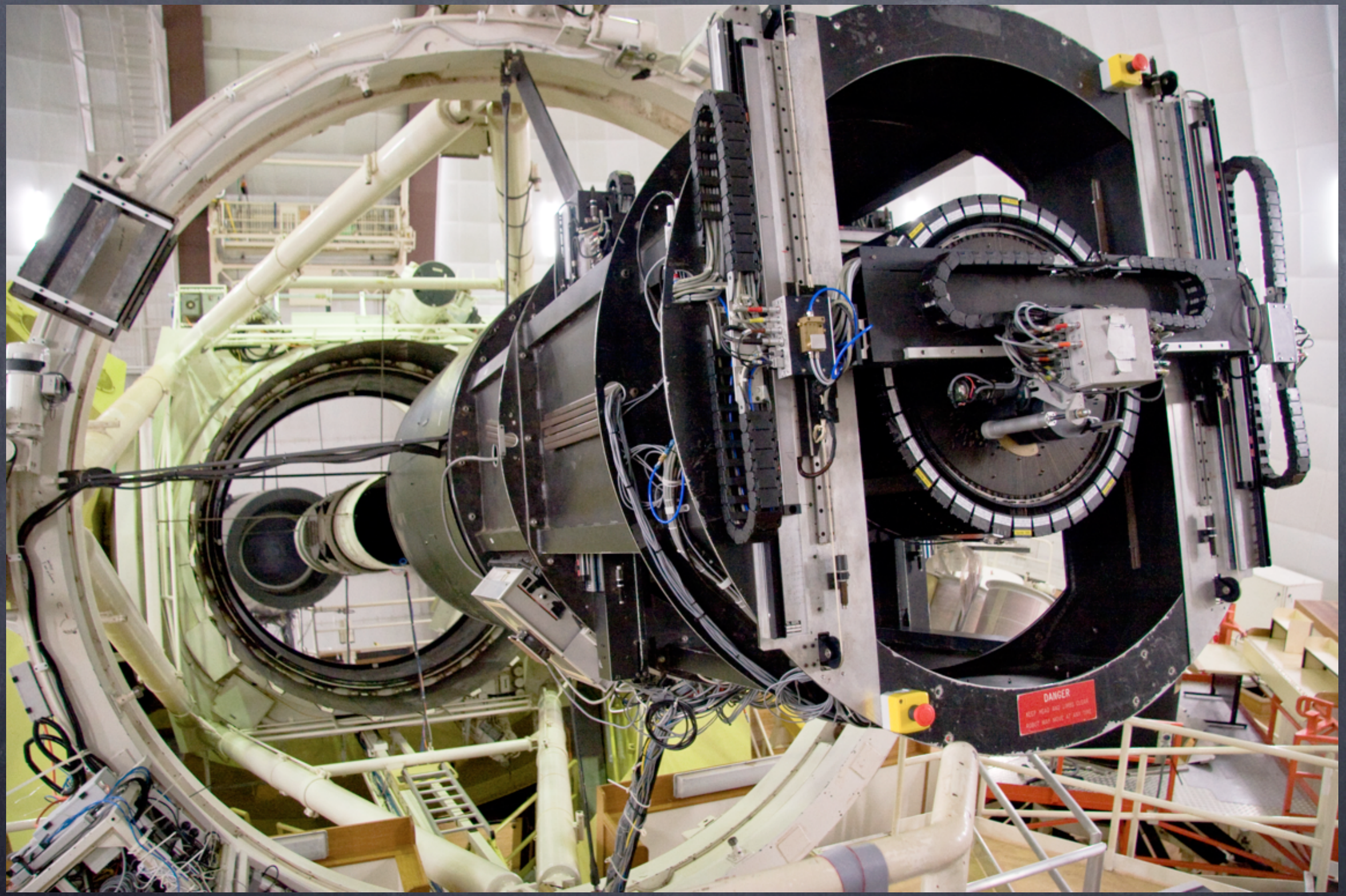
The GALAH data sets will yield a comprehensive view of the formation and evolution of the Milky Way disk and address the following basic questions:

- What were the conditions of star formation during early stages of Galaxy assembly?
- When and where were the major episodes of star formation in the disk and what drove them?
- To what extent is the Galactic disk composed of stars from merger events?
- Under what conditions and in what types of systems did accreted stars form?
- How have the stars that formed *in situ* in the disk evolved dynamically since their birth?
- Where are the solar siblings that formed together with our Sun?

Our current priority is small scale projects that will allow us to test the HERMES instrument.

Scientific Goals

After their birth-clusters or birth-galaxies disperse, stars may change their dynamical behavior thanks to mechanisms like heating and radial migration. However, the chemical composition of these stars, which reflect the conditions of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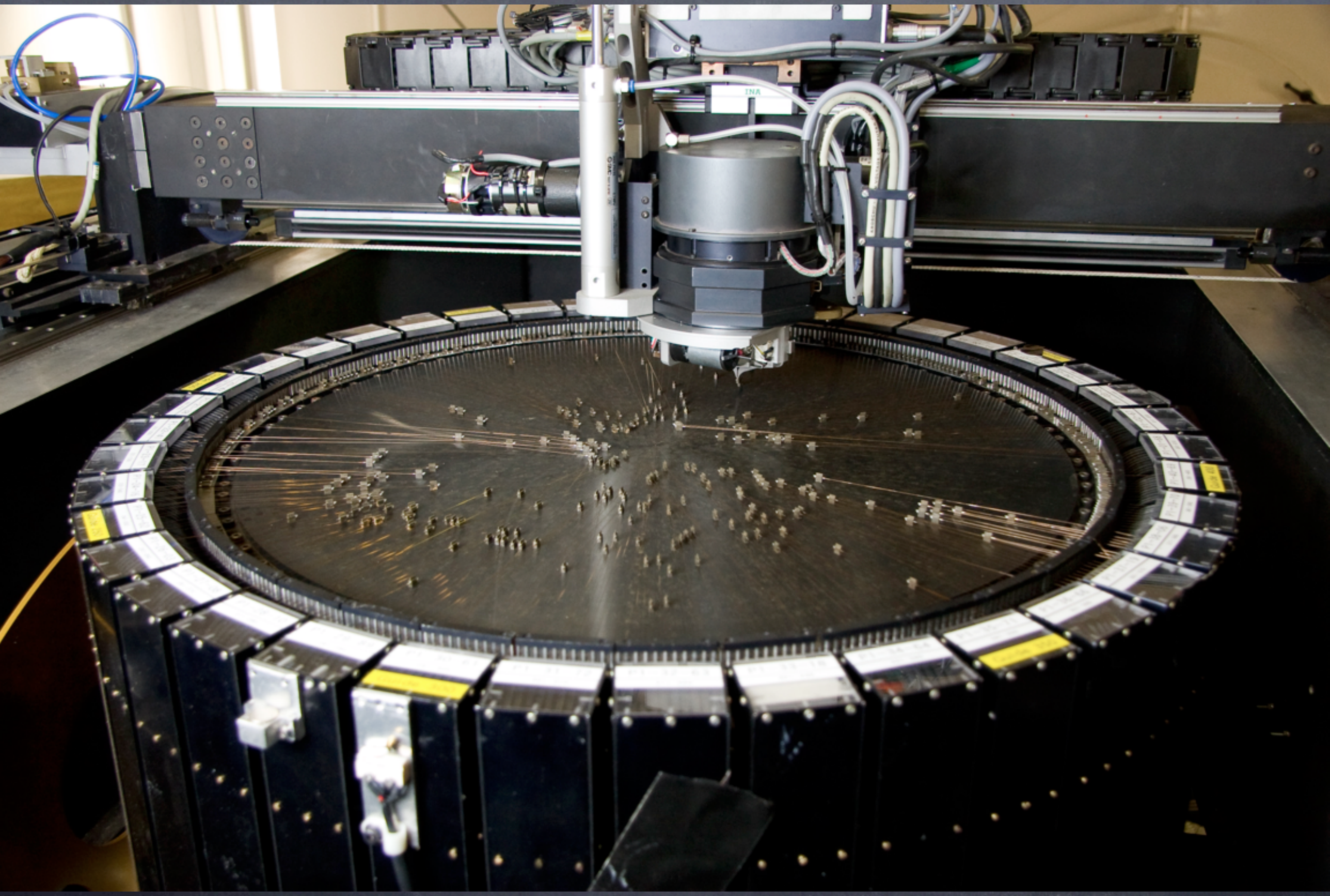
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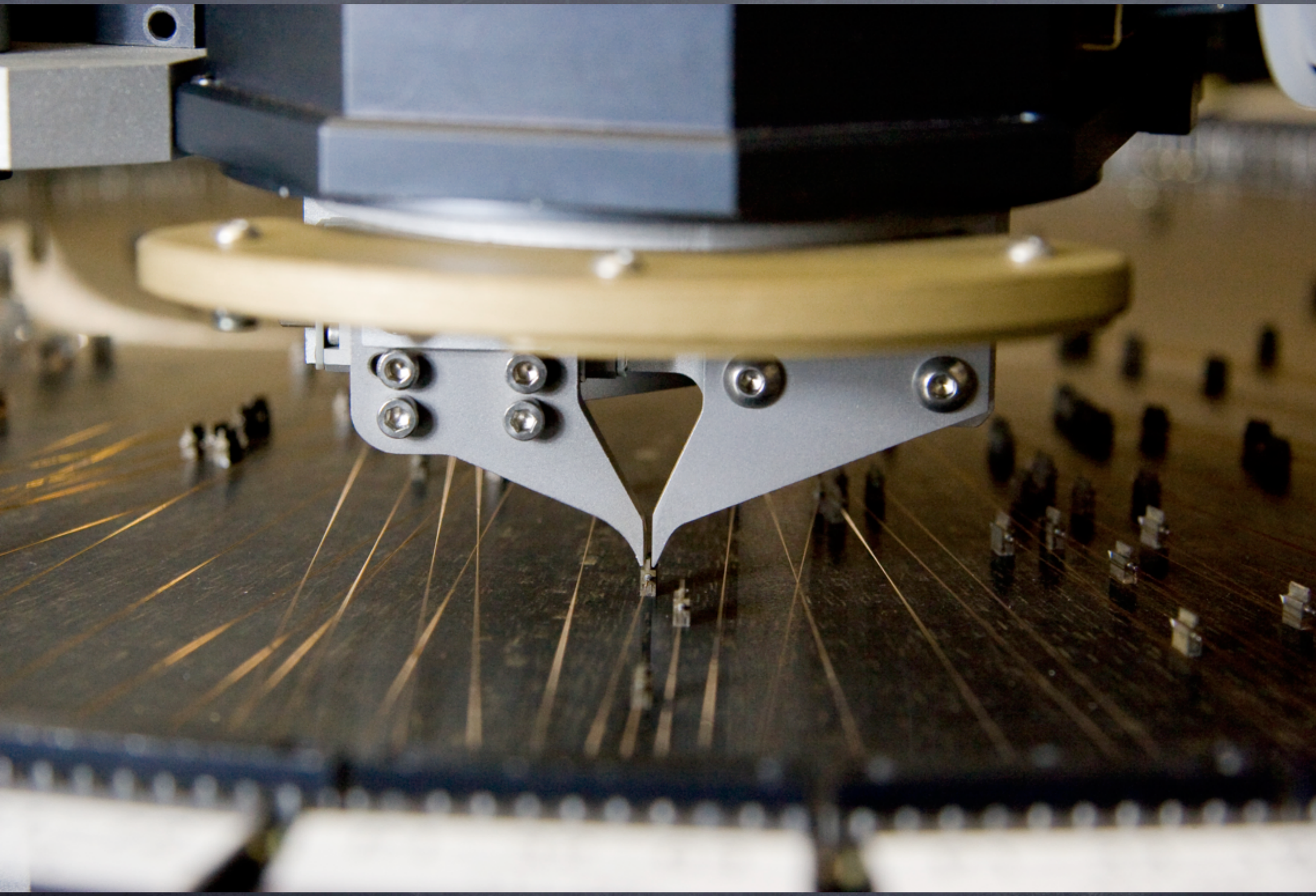

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AAOMEGA OBSERVATIONS OF 47 TUCANAE: EVIDENCE FOR A PAST MERGER?

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ABSTRACT

The globular cluster 47 Tucanae (47 Tuc) is well studied but it has many characteristics that are unexplained, including a significant rise in the velocity dispersion profile at large radii, indicating the exciting possibility of two distinct kinematic populations. In this Letter, we employ a Bayesian approach to the analysis of the largest available spectral data set of 47 Tuc to determine whether this apparently two-component population is real. Assuming the two models were equally likely before taking the data into account, we find that the evidence favors the two-component population model by a factor of $\sim 3 \times 10^7$. Several possible explanations for this result are explored, namely, the evaporation of low-mass stars, a hierarchical merger, extant remnants of two initially segregated populations, and multiple star formation epochs. We find the most compelling explanation for the two-component velocity distribution is that 47 Tuc formed as two separate populations arising from the same proto-cluster cloud which merged $\lesssim 7.3 \pm 1.5$ Gyr ago. This may also explain the extreme rotation, low mass-to-light ratio, and mixed stellar populations of this cluster.

Key words: globular clusters: individual (47 Tucanae)

Online-only material: color figure

1. INTRODUCTION

estimates, the membership selection process, and statistical analysis of cluster membership for all data presented in this

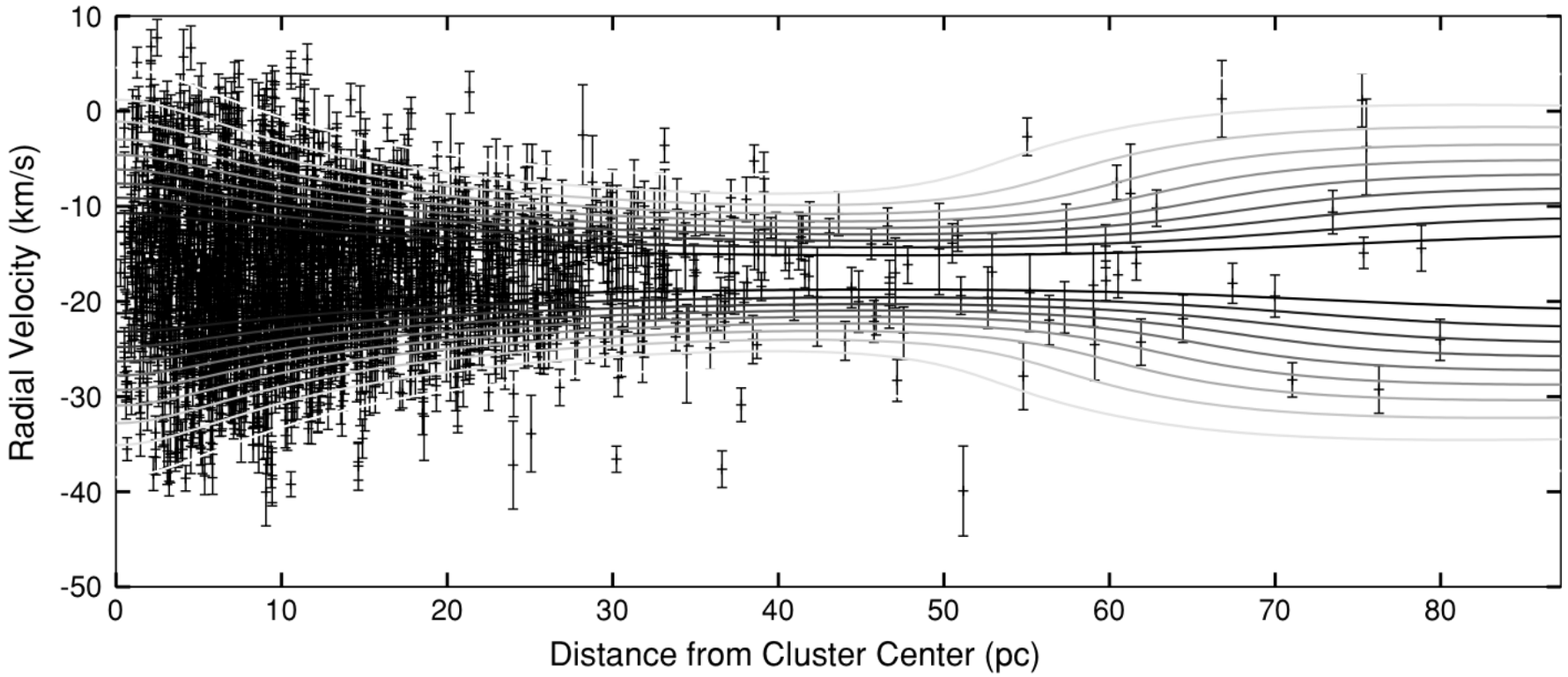

Figure 1. Radial velocities of stars in 47 Tuc together with the best-fit double Plummer model for the velocity distribution as a function of radius. The inner part of the cluster is well modeled by the dominant Plummer profile, while at larger radii, the second plummer profile dominates. The radius at which the two profiles have equal weight is 55 pc.

Table 1
Inferred Parameter Values for the Single Plummer and the Double Plummer Fits to the 47 Tuc Data

Parameter	Value
Single Plummer Profile	
μ	-16.87 ± 0.17 km s^{-1}
σ_0	9.37 ± 0.32 km s^{-1}
r_0	9.27 ± 0.98 pc
log(evidence)	-7759.5
Double Plummer Profile	
μ	-16.94 ± 0.12 km s^{-1}
σ_0	9.93 ± 0.43 km s^{-1}
	[5.70, 13.51] km s^{-1}

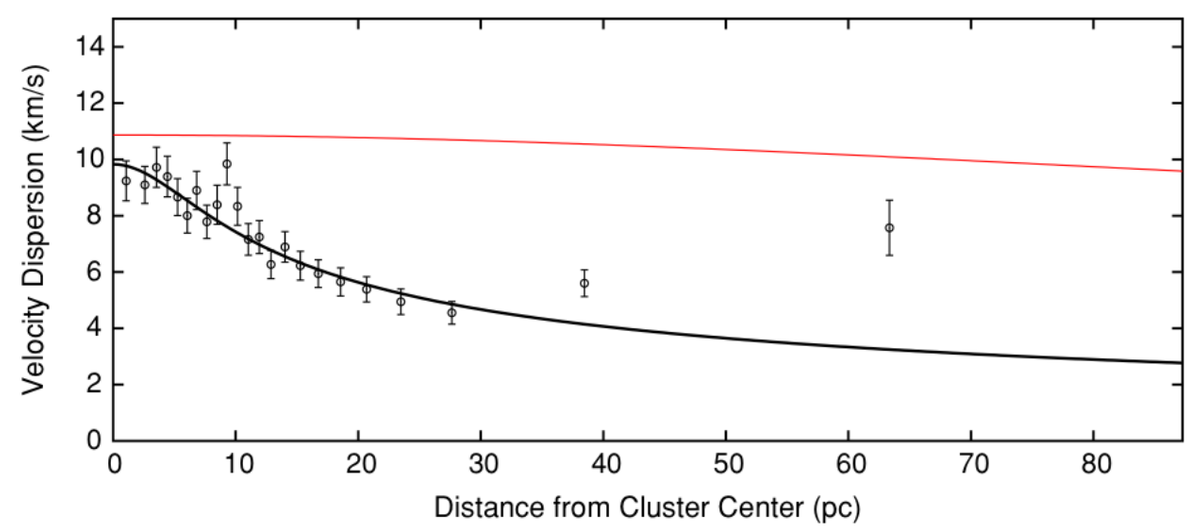

Figure 2. Binned velocity dispersion as a function of radius (from Lane et al. 2010a), with the radial velocity profiles of the two stellar populations from the best-fit double Plummer model. The Plummer profile that dominates at small radii is shown as the thick (black) curve; the thin (red) curve shows the Plummer

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WEAVE

A new wide-field multi-object spectrograph for the William Herschel Telescope

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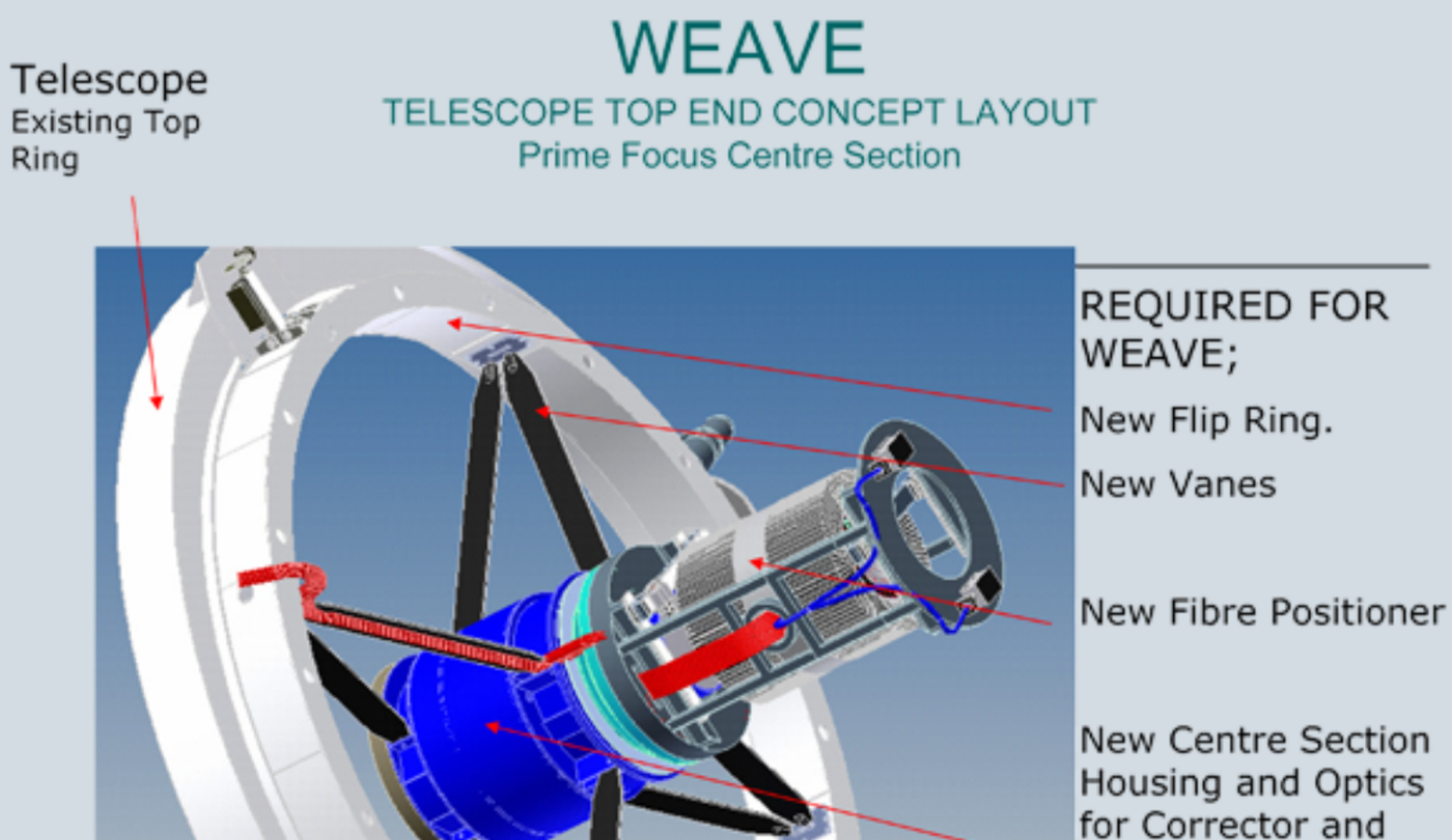

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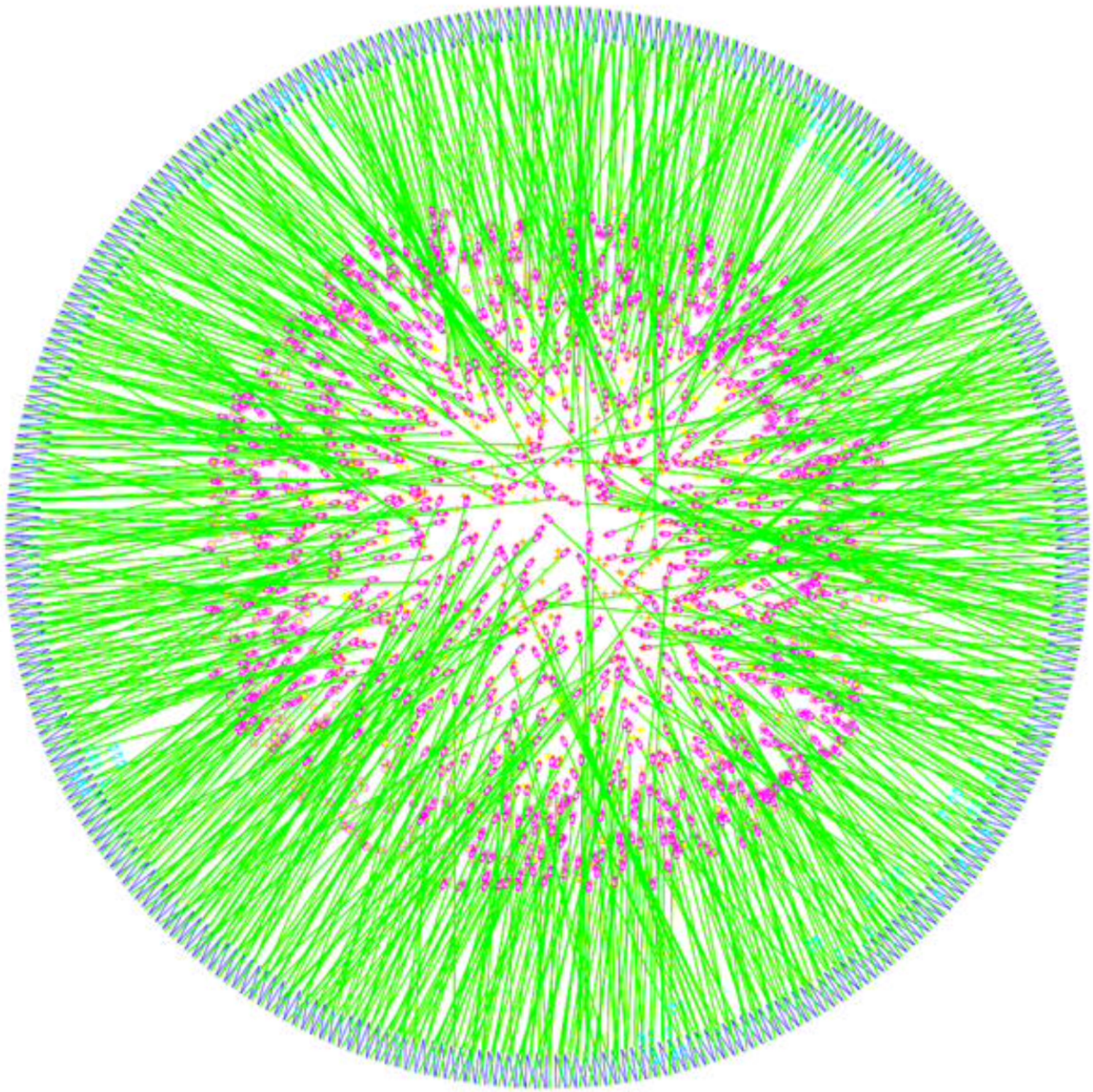
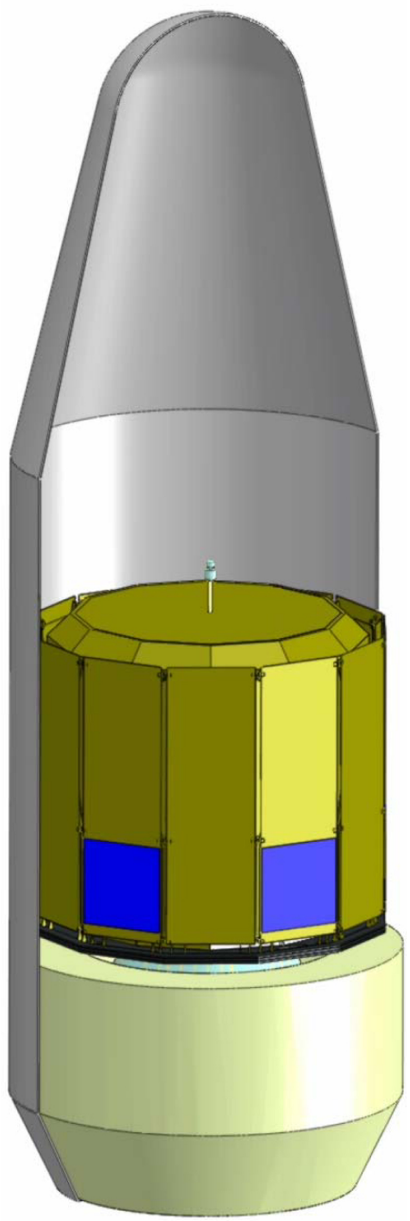
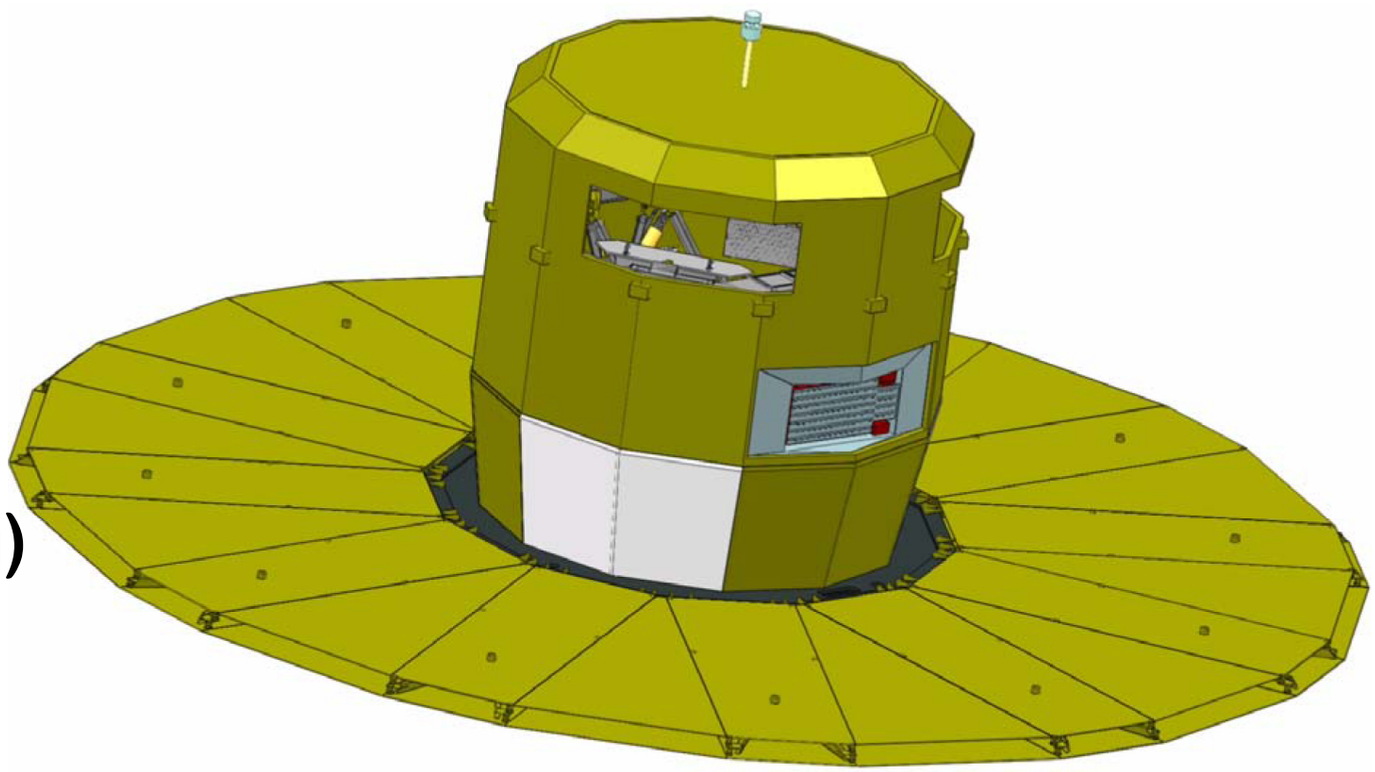


Gaia: asztrometriai űrobszervatórium



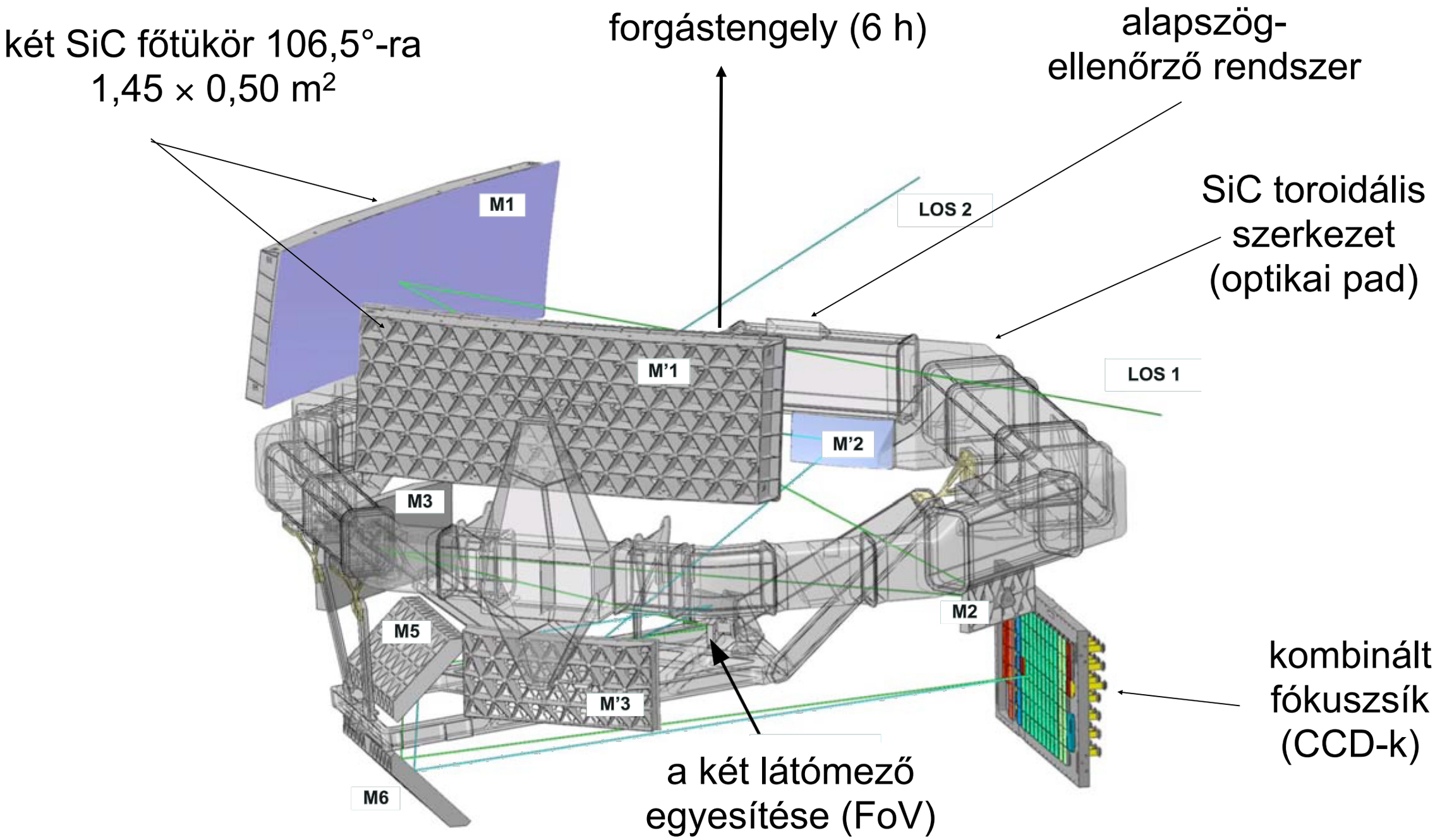
- kizárólag ESA-misszió
- indítás: 2013. december 20.
- 5 év működés (esetleg 1 év hosszabbítás)
- hordozórakéta: Szojuz–Fregat
- pálya: Lissajous-pálya az L2 pont körül
- adattovábbítás: 4–8 Mbps



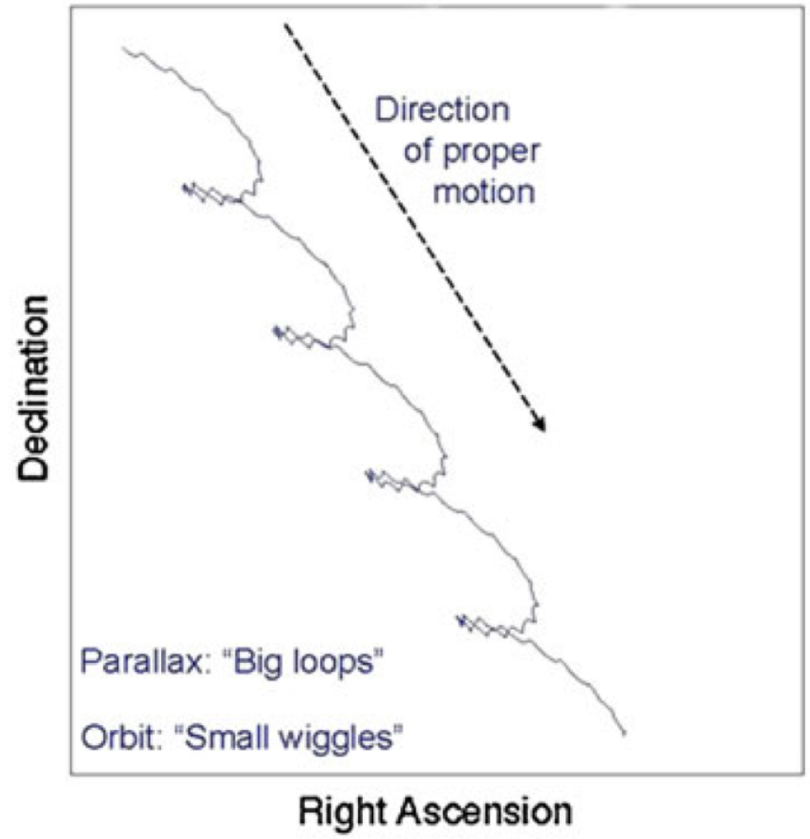
- tömeg: 2120 kg (hasznos 743 kg)
- teljesítmény: 1631 W (hasznos 815 W)

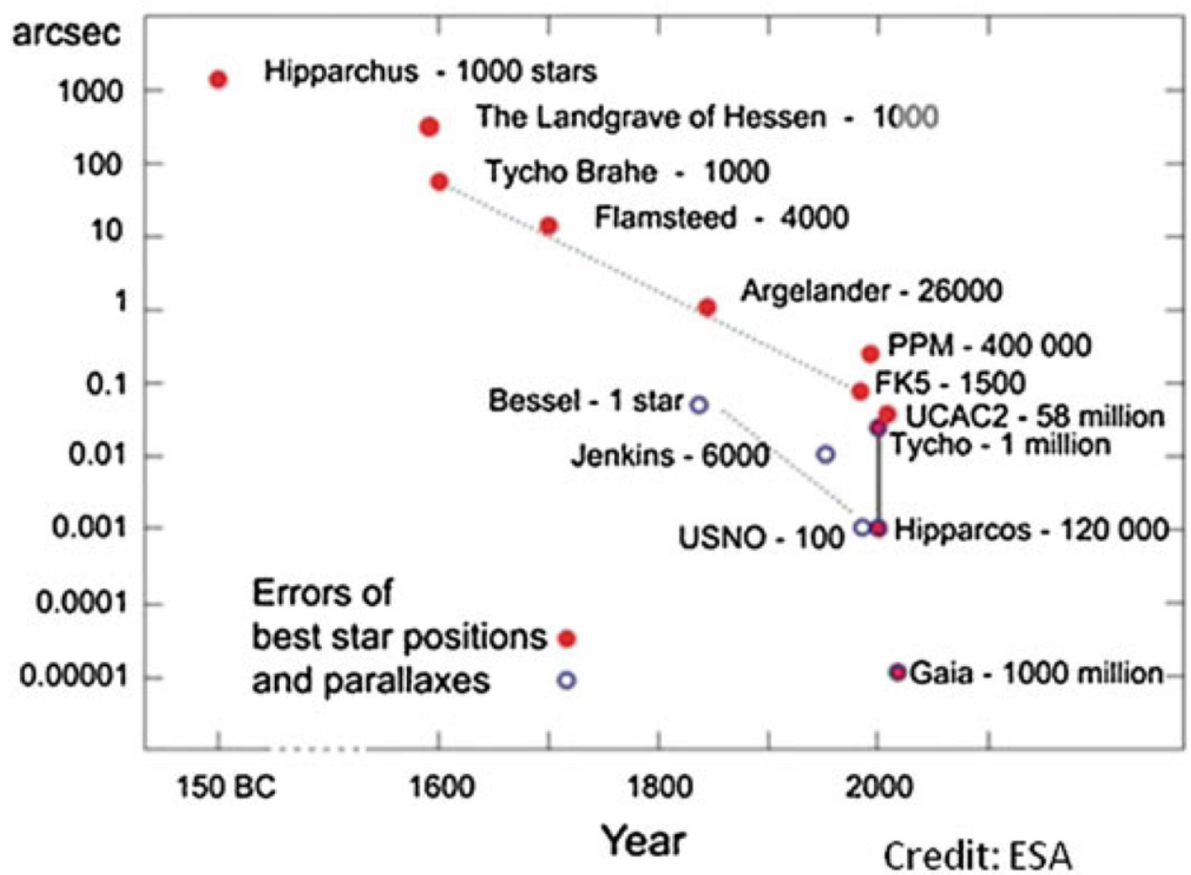
EADS-Astrium alapján

A távcső és a műszerek



EADS-Astrium alapján


Direction
of proper
motion
Declination
Parallax: "Big loops"
Orbit: "Small wiggles"
Right Ascension


arcsec
1000
100
10
1
0.1
0.01
0.001
0.0001
0.00001
Hipparchus - 1000 stars
The Landgrave of Hessen - 1000
Tycho Brahe - 1000
Flamsteed - 4000
Argelander - 26000
PPM - 400 000
FK5 - 1500
UCAC2 - 58 million
Tycho - 1 million
Bessel - 1 star
Jenkins - 6000
USNO - 100
Hipparcos - 120 000
Gaia - 1000 million
Errors of
best star positions
and parallaxes
150 BC
1600
1800
2000
Year
Credit: ESA

Gaia: teljes, pontos, mély

	Hipparcos	Gaia
Magnitúdóhatár	12	20 magnitúdó
Teljesség	7,3-9,0	20 magnitúdó
Fényes határ	0	6 magnitúdó
Objektumok száma	120 000	26 millió V = 15 m-ig,
		250 millió V = 18 m-ig
		1000 millió V = 20 m-ig
Effektív távolsághatár	1 kpc	50 kpc
Kvazárok	1 (3C 273)	500000
Galaxisok	nincs	1000000
Pontosság	0,001 ívmásodperc	7 mikroívmásodperc V = 10 m
		10-25 mikroívmásodperc V = 15 m
		300 mikroívmásodperc V = 20 m
Fotometria	2-színfotometria (B és V)	alacsony diszp. spektrum V = 20 m-ig
Radiális sebesség	nincs	15 km/s V=16-17 m-ig
Megfigyelési program	előre kiválasztva	teljes es torzítatlan

A pontosság 2 nagyságrendet javul, az érzékenység 4 nagyságrenddel jobb, a vizsgált csillagok száma 4 nagyságrenddel több.

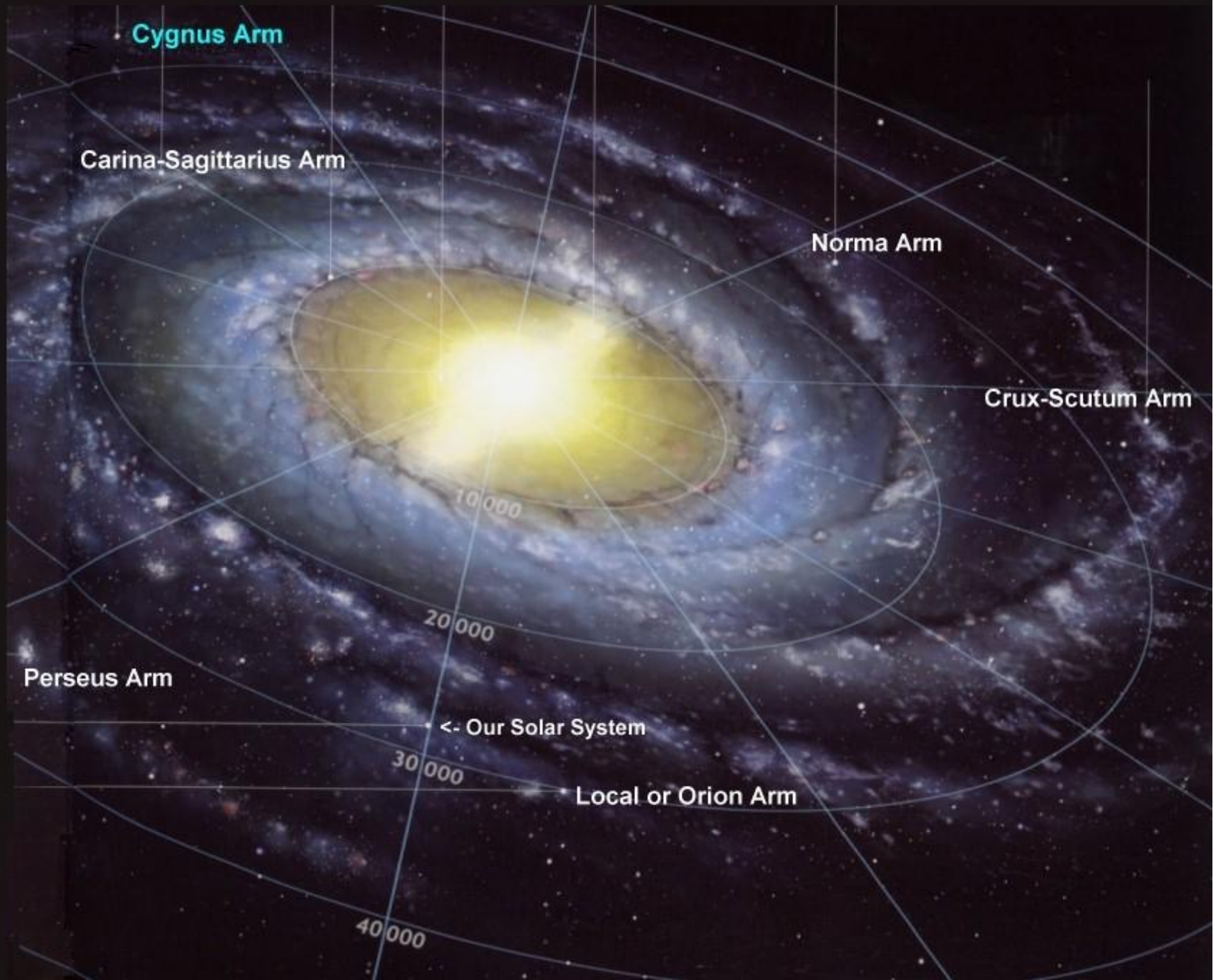
Az égbolt szkennelése 5 éven át ⇒ parallaxisok és sajátmozgások.

Gaia az asztrofizika számára

- A luminozitás átfogó kalibrálása, egyebek között:
  - 1% pontos távolságok ~10 millió csillagra 2,5 kpc-en belül
  - 10% pontos távolságok ~100 millió csillagra 25 kpc-en belül
  - Számos ritka típusú, ill. gyors fejlődési fázisban levő csillag
  - Minden távolságindikátor parallaxisának kalibrálása: pl. cefeidák, RR Lyrae típusú csillagok a Magellán-felhőkben is
- Csillagok fizikai jellemzői, egyebek között:
  - pontos Hertzsprung–Russell-diagram az egész Tejútrendszerre
  - a Naphoz közeli csillagok tömegfüggvénye és luminozitási függvénye, pl. fehér törpék (~200000), barna törpék (~50000)
  - csillagkeletkezési régiók tömeg- és luminozitási függvénye
  - fősorozat előtti csillagok luminozitási függvénye
  - minden színképtípus és csillagpopuláció detektálása és korának meghatározása
  - minden színképtípusú csillag változásának detektálása és jellemzése

1 milliárd csillag 3 dimenziós vizsgálata

- a Tejútrendszerre vonatkozóan…
  - minden csillagpopuláció távolsága és sebességeloszlása
  - a korong és a haló térbeli és dinamikai szerkezete
  - keletkezésének története
  - a galaktikus sötét anyag eloszlásának részletes feltérképezése
  - a csillagszerkezetre és -fejlődésre vonatkozó elméletek szelektálása megfigyelési alapokon


Cygnus Arm
Carina-Sagittarius Arm
Norma Arm
Crux-Scutum Arm
10 000
20 000
Perseus Arm
<- Our Solar System
30 000
Local or Orion Arm
40 000